

Empirical models of Total Electron Content based on functional fitting over Taiwan during geomagnetic quiet condition

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Abstract. Empirical models of Total Electron Content (TEC) based on functional fitting over Taiwan (120° E, 24° N) have been constructed using data of the Global Positioning System (GPS) from 1998 to 2007 during geomagnetically quiet condition ($D_{st} > -30$ nT). The models provide TEC as functions of local time (LT), day of year (DOY) and the solar activity (F), which are represented by 1–162 days mean of F10.7 and EUV. Other models based on median values have been also constructed and compared with the models based on the functional fitting. Under same values of F parameter, the models based on the functional fitting show better accuracy than those based on the median values in all cases. The functional fitting model using daily EUV is the most accurate with 9.2 TECu of root mean square error (RMS) than the 15-days running median with 10.4 TECu RMS and the model of International Reference Ionosphere 2007 (IRI2007) with 14.7 TECu RMS. IRI2007 overestimates TEC when the solar activity is low, and underestimates TEC when the solar activity is high. Though average of 81 days centered running mean of F10.7 and daily F10.7 is often used as indicator of EUV, our result suggests that average of F10.7 mean from 1 to 54 day prior and current day is better than the average of 81 days centered running mean for reproduction of TEC. This paper is for the first time comparing the median based model with the functional fitting model. Results indicate the functional fitting model yielding a better performance than the median based one. Meanwhile we find that the EUV radiation is essential to derive an optimal TEC.

Keywords. Ionosphere (Equatorial ionosphere; Ionospheric disturbances) – Radio science (Ionospheric physics)

1 Introduction

The dual frequency radio signals of the Global Positioning System (GPS) are widely used to estimate ionospheric Total Electron Content (TEC), which is one of the most powerful techniques to observe ionosphere. Many ionospheric phenomena, such as traveling ionospheric disturbances (Saito et al., 2001), large scale traveling ionospheric disturbances during magnetic storm (Tsugawa et al., 2003) and tsunami ionospheric disturbance after a large earthquake (Liu et al., 2006a), have been revealed by GPS-TEC. One of the most pronounced researches investigated by GPS-TEC is pre-seismic ionospheric perturbations (Liu et al., 2000, 2001, 2004). The ionospheric perturbations caused by the earthquakes are small as compared with local time (LT), seasonal and solar flux variation. Thus references for the researches are quite important to study the precursor phenomena related to the earthquakes. Previous studies used 15-days running median of GPS TEC, which is defined as previous 15 days median, as the reference for earthquake researches (Liu et al., 2000, 2001, 2004). Though the 15-days running median is very useful, the running median has a shortcoming during high solar activity because the 15-days running median cannot follow large fluctuation of the solar activity F10.7, whose variable periods are from 13 to 40 days (Kane, 2003; Uozumi et al., 2008). The most well known reference, International Reference Ionosphere (IRI), which has been developed since 1970s (Rawer et al., 1978), is very good tool to reproduce global disturbances of lower ionosphere. However, IRI has often overestimates plasma density in upper ionosphere and TEC as many authors have pointed out (Bilitza and Williamson, 2000; Iwamoto et al., 2002; Bilitza, 2001, 2004; Bilitza et al., 2006; Kakinami et al., 2008). Hence it is underqualified to use IRI as the reference such as small and localized variation of ionosphere associated



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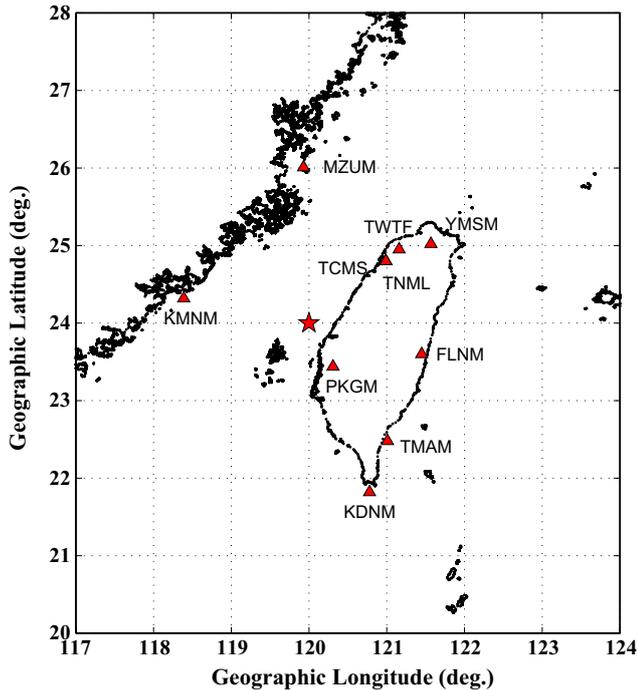


Fig. 1. GPS receivers in the Taiwan area. Red triangles and red star denote the locations of GPS receivers and the place where the TEC models are constructed (24°N 120°E), respectively.

with the earthquakes. Recently, this shortcoming of IRI has been improved in the version of 2007 (IRI2007) by applying “NeQuick” (Coisson et al., 2008a) which based on DGR “profiler” (Di Giovanni and Radicella, 1990). Compared to IRI2001, IRI2007-NeQuick option improved the estimate of TEC (Coisson et al., 2008b).

In this paper, we construct empirical models of TEC based on functional fitting over the Taiwan area with several parameters of the solar activity. We also construct other empirical models based on median values to compare accuracies of the models. While our models have been constructed to investigate the ionospheric perturbation associated with the large earthquakes, since a regional empirical model has not been constructed yet in the place where TEC variation is large due to ionospheric equatorial anomaly such as the Taiwan area, comparisons of our models with IRI2007 will contribute to improve IRI.

2 Total Electron Content using the Global Positioning System

The GPS provides an ideal tool for monitoring temporal and spatial variations of ionospheric electron densities. Many authors (Sardon et al., 1994; Leick, 1995; Liu et al., 1996) evaluated the ionospheric TEC with measurements of modulations on carrier phases and phase codes recorded by dual-frequency ($f_1=1575.42$ MHz and $f_2=1227.60$ MHz) GPS re-

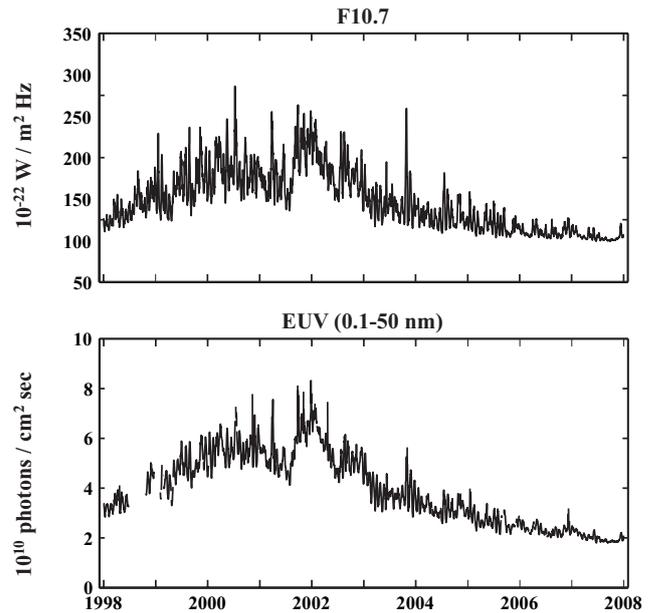


Fig. 2. Time series of F10.7 and EUV fluxes at 0.1–50 nm obtained by SOHO. EUV have been adjusted to the sun-earth distance. Spike-like sudden enhancements of EUV have been removed.

ceivers. However, instrumental biases are different for the dual-frequencies by each GPS receiver, which produces systematic errors in the estimates of TEC. In order to remove the instrumental biases, in this paper, the Global Ionospheric Map (GIM) developed by the Jet Propulsion Laboratory (Mannucci et al., 1998) is used as a reference to define the biases. The GIM TEC are routinely constructed and published in a 2-h interval by using GPS signals, which are obtained by 1000s of worldwide ground-based receivers. Note that the spatial resolutions of the GIM on the $\pm 87.5^\circ$ latitude and $\pm 180^\circ$ longitude are 2.5° and 5° , respectively. Therefore, we can obtain the instantaneous “snapshot” of the global TEC distribution in spatial and diurnal TEC variation at fixed location. According to Bilitza (2001), the minimum TEC in a day appears around 04:00–06:00 local time. Hence the differential instrumental biases of GPS receivers in the Taiwan area are determined by the minimum TEC value at 04:00–06:00 local time which are calculated by GIM at the place where GPS receivers are situated.

Combining all the data recorded by the GPS network in the Taiwan area (Fig. 1), the diurnal variation of TEC is examined, and then the latitude/longitude map of TEC over the Taiwan area is constructed. An elevation cut-off angle of each GPS receiver is 20 degrees. In order to investigate the daily TEC variations around the Taiwan area, the median value of TEC within 120° E and 23° N– 25° N, where the data are the densest, are calculated, and taken the 1.25 h running mean with 15 min time resolution to eliminate high frequent variations. The values of TEC from 1998 to 2007 are used to construct the models.

3 The solar activity

Variations of observed solar radio noise 10.7 cm, so-called F10.7, which are released on National Geophysical Data Center (<http://www.ngdc.noaa.gov/stp/SOLAR/ftpsolarradio.html>), is displayed in the top panel of Fig. 2. F10.7 gradually increased and showed the first peak around January 2000. F10.7 reduced toward the end of 2001, and then suddenly increased again and took the second peak around August 2001. After the second peak, F10.7 showed a continuous reduction toward the end of 2007. The range of F10.7 was between 65.0 and 314.6 during 1998–2007.

Daily EUV (0.1–50 nm) obtained by the Solar Heliospheric Observatory (SOHO) (Judge et al., 1998) are released on web site of Space Science Center in University of Southern California (<http://www.usc.edu/dept/space-science/>). Since the released EUV fluxes are adjusted to the flux at 1 AU, the released EUV fluxes are modified by using the sun-earth distance. Intensive spike-like enhancements in daily EUV, which exceeded almost one order of prior day’s intensity, were sometimes observed when many invasive flares occurred in one day. The modified EUV which are excluded sudden enhancements are shown in the bottom panel of Fig. 2. The variation of EUV is similar to F10.7 variation, which can be seen the double peaks during the solar maximum.

Here, indicators for the solar activity, F , are defined as follows:

$$F_{xdc} = (F107C + F107d)/2, \tag{1}$$

$$F_{xdp} = (F107P + F107d)/2, \tag{2}$$

$$F_{EUV} = EUVd, \tag{3}$$

where

$$F107C = \sum_{d=-(x-1)/2}^{d=(x-1)/2} F107d / x, \tag{4}$$

$$F107P = \sum_{d=-x}^{d=-1} F107d / x, \tag{5}$$

respectively. F107d and EUVd are the daily observed F10.7 and the daily modified EUV, respectively. X , which is a length of day, are 1, 7, 15, 27, 54, 81, 108, 135 and 162 days. Namely, F107C and F107P represents centered F10.7 mean of X days and F10.7 mean from 1 to X days before, respectively. SOHO observed two EUV bands, which are 26–34 nm and 0.1–50 nm. These two bands are correlated linearly (Bilitza, 2000; Kane, 2003; Liu et al., 2006b). The most effective EUV band for TEC production is 65–91 nm, which produces O^+ in F1 region (Banks and Kockarts, 1973). From comparison of 0.1–50 nm band intensity with 65–90 nm band intensity during 1994–2007 reproduced by a solar EUV model (Richards et al., 1994a, b), intensities of these two bands are

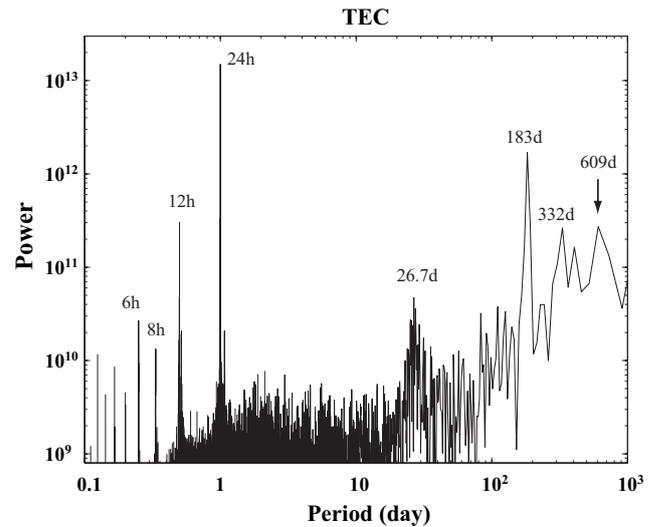


Fig. 3. Power spectrum of TEC variation over the Taiwan area from 1998 to 2007.

highly correlated (correlation coefficient is 0.99). Thus, 0.1–50 nm band is reasonable for the parameter of TEC reproduction. The band of EUV flux in 0.1–50 nm is only used hereafter.

In order to investigate coherence of F_{EUV} to F_{xdc} and F_{xdp} , linear correlation coefficients, $S(x, y) = C(x, y) / \sqrt{C(x, x)C(y, y)}$, where the covariance between data set x and y , $C(x, y) = \mu[(x - \mu_x)(y - \mu_y)]$ and μ_x is the mathematical expectation of x , are examined. Though F_{81dc} has been pointed out as good indicator for EUV (Liu et al., 2006b), correlation between F_{81dc} and F_{EUV} does not show the strongest value with 0.981 (see Table 1). The correlation coefficients for F_{54dp} , F_{81dp} , F_{108dp} , F_{108dc} , F_{135dc} and F_{162dc} are slightly better than that of F_{81dc} , which are 0.982.

4 TEC variation over the Taiwan area

To determine the functions for the models based on the functional fitting, TEC variations are investigated before the functions are selected. In following subsection, diurnal, seasonal, annual and longer variations are discussed. In addition, responses of TEC with the solar activity are also shown.

4.1 Diurnal variation

Figure 3 shows power spectrum of TEC from 1998 to 2007. Periods of 4, 8, 12, and 24 h are clearly seen. Since ionospheric plasma is mainly produced by EUV from the sun, the period of 24 h is the most dominant. Other periods might come from interaction between neutral atmosphere and ionosphere. Recently, evidences of modulation caused by waves of neutral atmosphere have been found (Sagawa et al., 2005;

Table 1. The correlation coefficients of $F_{x dp}$ and $F_{x dc}$ with EUV (0.1–50 nm).

	1d	7d	15d	27d	54d	81d	108d	135d	162d
$F_{x dp}$	0.957	0.966	0.972	0.979	0.982	0.982	0.982	0.981	0.980
$F_{x dc}$	0.954	0.960	0.968	0.973	0.978	0.981	0.982	0.982	0.982

Table 2. The correlation coefficients between F parameters and TEC in 15:00 LT. The correlation coefficient between F_{EUV} and the TEC in 15:00 LT is 0.816.

	1d	7d	15d	27d	54d	81d	108d	135d	162d
$F_{x dp}$	0.776	0.791	0.801	0.804	0.809	0.809	0.809	0.809	0.808
$F_{x dc}$	0.773	0.780	0.791	0.801	0.807	0.809	0.809	0.809	0.808

Immel et al., 2006; Lin et al., 2007). For instance, eastward tide of wave number 3 is strongly affected equatorial ionosphere via collision between neutral atmosphere and plasma in E-layer. One of the phenomena caused by the interaction between neutral atmosphere and plasma is found in ionospheric EUV emission obtained by the IMAGE satellite (Sagawa et al., 2005; Immel et al., 2006). In addition, similar longitudinal structure of ionospheric plasma density was observed by FORMSAT3/COSMIC (Lin et al., 2007).

4.2 Seasonal, annual and longer variations

Since the rotation period of the sun is about 25 days, the synodic period with the earth, which produces F10.7 variation, is about 27 days. Therefore, the power spectrum has peak power in 26.7 days period. In F region, where peak plasma density is observed, production rate of plasma is dominated by density of O due to direct photo ionization of O and recombination rate is dominated by density of N_2 due to recombination of O^+ and N_2 . The density variations of O and N_2 , which have two maximums per year in March and September, produce the semiannual variation of TEC, whose period is 183 days. Density of F region is also affected by neutral wind in thermosphere (Oyama et al., 1996). The seasonal variation of the neutral wind also produces the seasonal variation of TEC. The power spectrum shows annual (332 day) and longer variations (609 day). Existence of the equatorial quasi-biennial oscillation (QBO) in stratosphere, whose period is approximately 28 months, is widely known (Baldwin et al., 2001). Since a troposphere-thermosphere coupling has been pointed out (Forbes et al., 2006), QBO might affect variation of plasma in the ionosphere. As the result of the coupling, longer variation of TEC whose period is 609 days might be produced. Coupling between ionospheric plasma and QBO will be studied in near future.

4.3 Response of TEC to the solar activity indexes, F

In order to study responses of TEC to the solar activity indexes F which was defined in Sect. 3, the correlation coefficients between F and TEC in 15:00 LT (TEC_{15LT}), which usually show daily maximum, during geomagnetic condition ($D_{st} > -30$ nT) are examined (Table 2). The correlation coefficient between F_{EUV} and the TEC_{15LT} shows the strongest value with 0.816. This result is no wonder because EUV is a primary source for ionospheric plasma above D-region (Bilitza, 2000) and F1 region (Banks and Kockarts, 1973). $F_{x dp}$ are better than $F_{x dc}$ for indicator of TEC_{15LT} when the same day lengths are compared. Although the sudden enhancements of EUV have been excluded in our analysis (the bottom panel of Fig. 2), if all EUV data are used for the analysis, a correlation coefficient between EUV and TEC_{15LT} (0.645) becomes weaker than the other coefficients. This fact suggests that TEC does not enhance linearly anymore when large flares frequently occurred. Seasonal variations of TEC_{15LT} with F_{1dc} parameter are illustrated in Fig. 4. The variations of TEC_{15LT} are saturated over F10.7=200 except September equinox. Clear saturation of TEC_{15LT} cannot be seen during September equinox. TEC_{15LT} over F10.7=250 in September equinox were observed September and October 2001. More than 10 geomagnetic storms occurred during this period. Large-scale equatorward moving gravity waves induced by Joule heating of auroral electro jet propagate during the geomagnetic storm (Richmond, 1978; Fujiwara et al., 1996). The equatorward wind arising from the gravity waves lifts plasma up in middle latitude, and then TEC enhance during the geomagnetic storm (Shiokawa et al., 2002). Although the data were selected during $D_{st} > -30$ nT, the frequent geomagnetic storms might affect ionosphere distribution longer than usual. Therefore TEC_{15LT} show such high values and the saturation might not be seen in September equinox in our study.

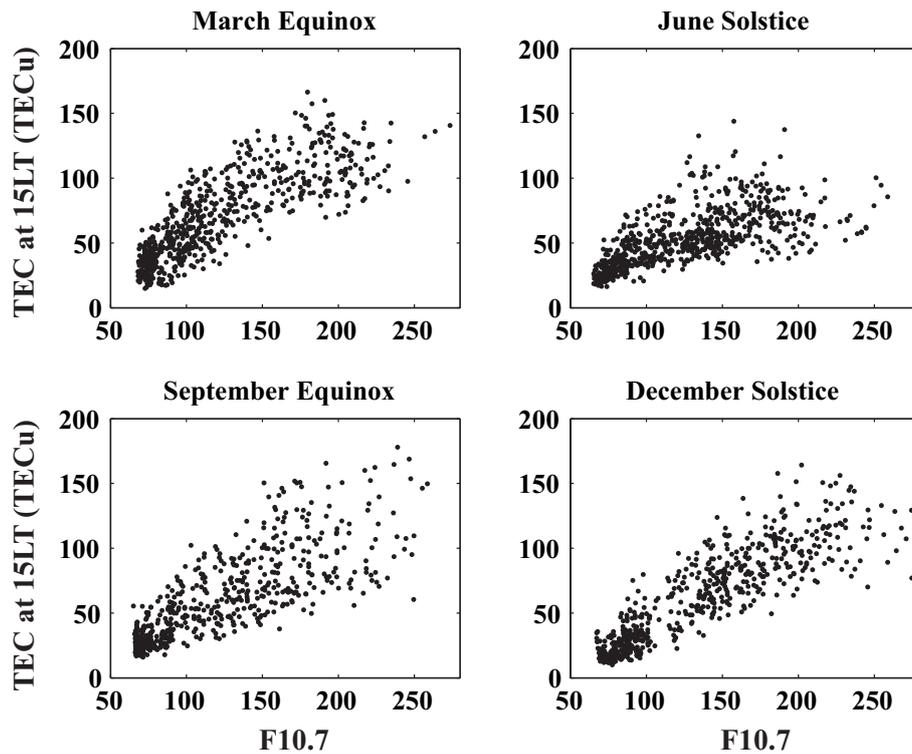


Fig. 4. TEC variation of 15:00 LT with daily F10.7 during the geomagnetic quiet condition ($D_{st} > -30$).

5 Methodology

Methodology of constructing empirical models based on the functional fitting is shown in following subsection. For comparison, we also construct other empirical models based on the median values, which is also described in following subsection. Both of the models are calculated TEC as functions for 3 parameters which are local time (LT), day of year (DOY) and the solar activity parameters F which are defined in the Sect. 4. TEC observed over the Taiwan area (120° E, 24° N) from 1998 to 2007 during the geomagnetic quiet condition ($D_{st} > -30$) are used to construct the both models. These data are not included whole of 11 year’s solar cycle but include the solar maximum and minimum conditions (Fig. 2).

5.1 The empirical model based on the functional fitting

The methodology of constructing empirical models based on the functional fitting, which has been originally applied and succeed to construct an empirical model of $O^+ - H^+$ transition height based on topside sounder data (Marinov et al., 2004), shown here. As shown in the previous section, since the variation of TEC with the solar activity show the saturation, cubic function or more high order function should be applied as the function for F parameters to reproduce the variation. Wave number 2 is required for seasonal (DOY) variation at least and wave number 4 is required for LT variation. As the

result of trial and error, the cubic function for F parameters, wave number 3 for the seasonal variation and wave number 4 for LT variation are applied to the functional models. The functions for each variation are defined as follows:

$$f(F) = a_1 + a_2F + a_3F^2 + a_4F^3 = a_1f'_1 + a_2f'_2 + a_3f'_3 + a_4f'_4, \tag{6}$$

$$g(\text{DOY}) = b_1 + \sum_{i=1}^3 (b_{2i} \cos 2\pi i \cdot \text{DOY} + b_{2i+1} \sin 2\pi i \cdot \text{DOY}) = b_1g'_1 + \dots + b_7g'_7, \tag{7}$$

$$h(\text{LT}) = c_1 + \sum_{i=1}^4 \left(c_{2i} \cos \frac{\pi i \cdot \text{LT}}{12} + c_{2i+1} \sin \frac{\pi i \cdot \text{LT}}{12} \right) = c_1h'_1 + \dots + c_9h'_9, \tag{8}$$

where a, b, c are coefficient for fitting and detail description of functions $f', g',$ and h' are displayed in Appendix A. DOY is normalized by the length of the year. Function reproducing TEC value is assumed as multiplication of these 3 functions. Then we can obtain:

$$\begin{aligned} \text{TEC}(F, \text{DOY}, \text{LT}) &= f(F) \cdot g(\text{DOY}) \cdot h(\text{LT}) \\ &= \sum_{i,j,k} a_i f'_i \cdot b_j g'_j \cdot c_k h'_k \\ &= \sum_{n=1}^{252} \alpha_n \text{TEC}'_n, \end{aligned} \tag{9}$$

Table 3. Bin ranges for F parameter: $A=[60, 80, 120, 160, 210, 300]$, $B=[60, 80, 120, 150, 200, 240]$; EUV=[1.8, 2.3, 3, 4, 5, 6.5] ($\times 10^{10}$ photons/cm² s).

Model	1d	7d	17d	27d	54d	81d	108d	135d	162d
Centered F10.7	A	A	A	A	B	B	B	B	B
Previous F10.7	A	A	A	B	B	B	B	B	B

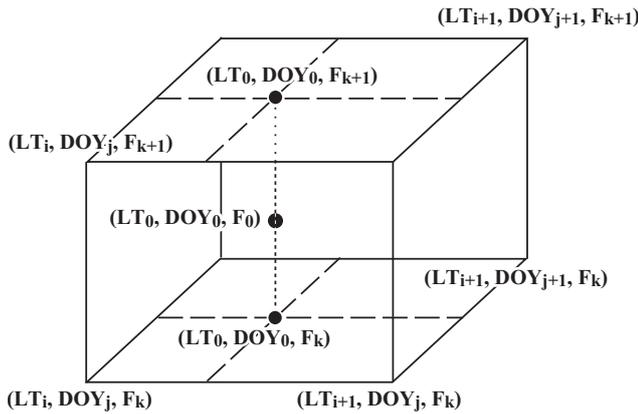


Fig. 5. The methodology of constructing the median models. LT_i , DOY_j , F_k denote the location of local time (LT), day of year (DOY), the solar activity (F) in (i, j, k) -th bin. The model values are calculated by cubic spline interpolation in LT-DOY space and by linear interpolation in F space.

where $n=(i-1)\times 63+(j-1)\times 9+k$, $i=1, \dots, 4$, $j=1, \dots, 7$, $k=1, \dots, 9$. In order to calculate coefficient α , more than 252 data are needed. The α can be obtained by ciphering following normal-equation matrix:

$$\begin{pmatrix} \sum \text{TEC} \cdot \text{TEC}'_1 \\ \vdots \\ \sum \text{TEC} \cdot \text{TEC}'_{252} \end{pmatrix} = \begin{pmatrix} \sum \text{TEC}'_1 \cdot \text{TEC}'_1 & \dots & \sum \text{TEC}'_1 \cdot \text{TEC}'_{252} \\ \vdots & \ddots & \vdots \\ \sum \text{TEC}'_{252} \cdot \text{TEC}'_1 & \dots & \sum \text{TEC}'_{252} \cdot \text{TEC}'_{252} \end{pmatrix} \begin{pmatrix} \alpha_1 \\ \vdots \\ \alpha_{252} \end{pmatrix}. \tag{10}$$

Several F are applied to develop the models, F_{xdc} functional model, $F_{x dp}$ functional model and F_{EUUV} function model. The values of coefficients α for F_{EUUV} functional model, which is the most accurate model as discussed later, are shown in Appendix B.

5.2 The empirical model based on the median values

The methodology of constructing the empirical models based on the median values is described here. In order to investigate response of TEC to the solar activities, several F parameters are applied to develop the models, F_{xdc} median model,

$F_{x dp}$ median model and F_{EUUV} median model. Three steps are taken to construct the models. First of all, bins are set in LT with 1 h resolution (24 nodes), DOY normalized by a length of the year with 1/12 resolution (12 nodes) and F space (5 nodes). The ranges of F are chosen as displayed in Table 3. The median values are calculated in each bin (LT_i, DOY_j, F_k) , where LT_i, DOY_j, F_k denote the location of LT, DOY, F in (i, j, k) -th bin and all median values are stored in a file. In second step, to calculate TEC at (LT_0, DOY_0, F_0) , TEC at (LT_0, DOY_0, F_k) and (LT_0, DOY_0, F_{k+1}) are interpolated using the cubic spline function in LT-DOY space from surrounding bins, where $LT_i < LT_0 < LT_{i+1}$, $DOY_j < DOY_0 < DOY_{j+1}$ and $F_k < F_0 < F_{k+1}$, respectively (also see Fig. 5). Finally, TEC at (LT_0, DOY_0, F_0) is interpolated by linear function in F space with (LT_0, DOY_0, F_k) and (LT_0, DOY_0, F_{k+1}) .

6 Model results and discussions

6.1 Time series of observed TEC, the 15-days running median, the function model and IRI2007

The TEC reproduced by the F_{EUUV} functional model are compared with the observed TEC, the 15-days running median and IRI2007 together with Dst and daily EUV (Fig. 6). During August 2002, two magnetic storms occurred around 2nd and 19th. Minimum Dst of the both storms were -96 nT. Daily EUV shows periodic variation, which increase from local minimal values of 2.2×10^{10} on 8th to local maximal value of 3.0×10^{10} on 19th and decrease to 2.3×10^{10} photos/cm² s on 27th. IRI2007 are in good agreement with the observed TEC from 23rd to the end of August, but obviously overestimate the observed TEC from 2nd to 12th. On the other hand, although the 15-days running median are in good agreement with the observed TEC from 2nd to 12th, the 15-days running median underestimate the observed TEC from 13th to 23rd. These disagreements come from rapid variation of daily EUV during 8th to 27th. The 15-days running median cannot clearly follow this kind of rapid F10.7 variation. In contrast, the F_{EUUV} function model is in very good agreement with the observed TEC except periods of the geomagnetic storm. Since our models are constructed by the data without the geomagnetic storms, our models only reproduce TEC during the quiet geomagnetic

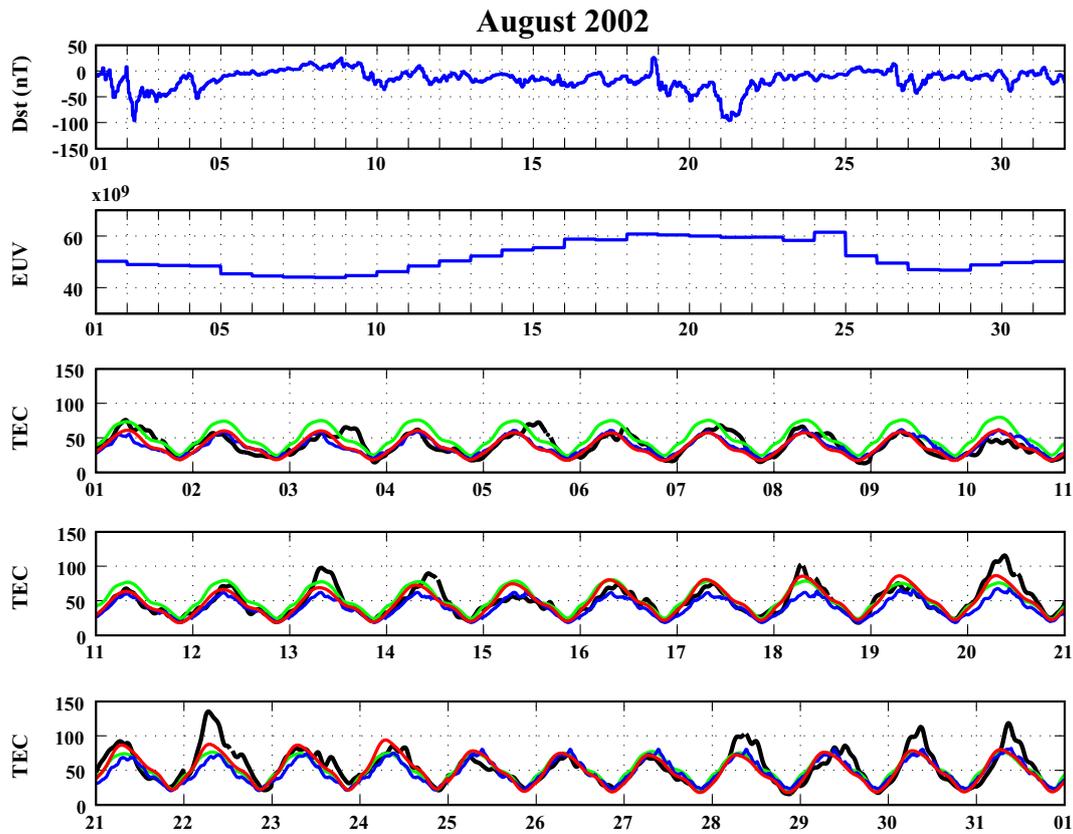


Fig. 6. Time series of Dst in 1st panel, daily EUV (1–50 nm) in 2nd panel, the observed TEC (black), the F_{EUV} function model (red), the 15-days running median (blue) and IRI2007 (green) in the bottom 3 panels during August 2002.

condition. In other ward, our models are good references for comparing the quiet to the disturbed ionosphere. Thus the enhancement of the observed TEC around 20th indicates an effect of the magnetic storm by the comparison of the observed TEC with the F_{EUV} functional model. When duration of the geomagnetic storm is long enough, the 15-days running median is contaminated by the magnetic storm effect itself. In such case, the 15-days running median cannot be referred as quiet ionosphere any more. This is another shortcoming of the 15-days running median.

6.2 Root mean square errors of the models

In order to estimate accuracies of the models, root mean square error (RMS), which is widely used to evaluate the model error (Kutiev and Marinov, 2007; Habarulema et al., 2007), is calculated as follows:

$$RMS = \sqrt{\frac{1}{n} \sum_{i=1}^n (o_i - m_i)^2}, \quad (11)$$

where o_i , m_i and n denote the observed TEC, the modeled TEC and total number of the data, respectively. Figure 7 illustrates RMSs for the 15-days running median (green triangle), the functional models (inverted triangle) and the me-

dian models (square) during the geomagnetic quiet condition, respectively. Colors of each mark display difference of the solar flux indicators; blue, red and yellow denote $F_{x dp}$, $F_{x dc}$ and F_{EUV} , respectively. Lengths of X days are shown in the horizontal axis. RMS of IRI2007 is 14.3 TECu, which is the biggest RMS value among all models. F_{EUV} is the most competent parameter for the empirical models. The F_{EUV} functional model shows the lowest RMS with 9.15 TECu. The F_{EUV} median model also shows the lowest RMS among the median models. These results are consequential because EUV radiations are the primary source for ionospheric plasma above D-region (Bilitza, 2000) and F1 region (Banks and Kockarts, 1973). Comparing with the same X days of F, the functional models show better accuracy than those of the median models. Though RMSs for the both two type of the models using $F_{1 dp}$ and $F_{1 dc}$ are similar 15-days running median, RMSs gradually decrease as the length of X days increase. The $F_{54 dp}$ functional model exhibits the minimum RMS with 9.38 TECu among the all models using F10.7. The $F_{54 dp}$ median model presents the best accuracy with 9.43 TECu among median models using F10.7. When longer X days are applied, accuracies of the model using F10.7 get worse again.

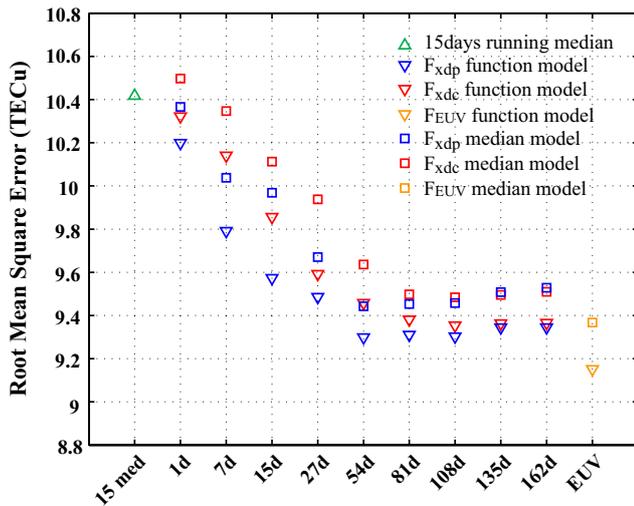


Fig. 7. Root mean square errors for the models. Marks show difference of the models types and colors show difference of F parameters. Namely, the 15-days running median (green triangle), the F_{xdp} functional model (blue inverted triangle), the F_{xdc} functional model (red inverted triangle), the F_{EUV} functional model (yellow inverted triangle), the F_{xdp} median model (blue square), the F_{xdc} median model (red square), and F_{EUV} median model (yellow square), respectively. Labels of X axis indicate the average lengths of F10.7.

6.3 Root mean square errors of the models for each year

RMSs are examined year by year to investigate difference of accuracies related to deference of the solar activity (Fig. 8). RMSs of the F_{EUV} functional model are more preferable than the 15-days running median and IRI2007 in all years. RMSs for the F_{EUV} function model and the 15-days running median become larger during the high solar activity and small during the low solar activity because RMS roughly correlates with the intensity of TEC. Differences of RMSs between the F_{EUV} functional model and the 15-days running median during the high solar activity are much larger than during the low solar activity. Response to the solar activity variation of the 15-days running median is much worse than that of the F_{EUV} functional model as shown in Fig. 7. Hence RMSs of the 15-days running median get worse during the high solar activity, which highly fluctuate, than the low solar activity. On the other hand, RMSs of IRI2007 are much larger than the others during the both high and low solar activity. IRI2007 can only reproduce TEC equivalently to the others during 1998–1999 and 2004–2005, which is the moderate solar activity. The reason why RMSs are large during the low and high solar activity arises from reproducibility of the solar flux variation as discussed later (the bottom panel of Fig. 10).

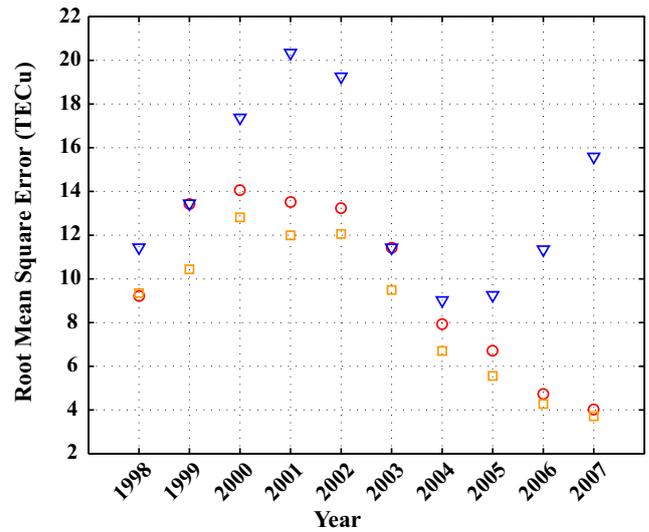


Fig. 8. Root mean square errors of IRI2007 (blue inverted triangle), the 15-days running median (red circle) and the F_{EUV} functional model (orange square) for each year from 1998 to 2007.

6.4 Comparison of the F_{EUV} functional model with IRI2007

LT, DOY, and the solar flux variations of the F_{EUV} functional model are investigated (Fig. 9). The medians of the ratio are between 0.96 and 1.01 in LT variation (the top panel of Fig. 9). The minimum and maximum median appear in 05:00 and 03:00 LT, where observed TEC are smaller than the daytime. Errors in the night time are also larger than that in the daytime. The F_{EUV} median model also shows good agreement with observations in DOY variation (middle panel of Fig. 9). The medians of the ratio are seen between 0.97 and 1.01. When F_{EUV} is over 7.5×10^{10} photon/cm² s, the F_{EUV} functional model slightly overestimates the observed TEC. The median of the ratio is 0.93.

Meanwhile, IRI2007 has large deviations to the observed TEC (Fig. 10). IRI2007 overestimates the observed TEC except 04:00 LT as illustrated in the top panel of Fig. 11. The medians of the ratio in LT variation are between 0.77 and 1.02. Regarding DOY variation, IRI2007 overestimates the observed TEC around June solstice. Particularly, IRI2007 largely overestimates the observed TEC in July. The median of the ratio is 0.80 in June. In addition, when F_{EUV} is below 3.0×10^{10} photon/cm² s, IRI2007 overestimates the observed TEC with decreasing F_{EUV} . Moreover, IRI2007 underestimates when F_{EUV} is over 6.0×10^{10} photon/cm² s. The median of the ratio reaches 1.2 in the case that F is 7.0×10^{10} photon/cm² s.

Comparing Figs. 9 to 10, the F_{EUV} functional model obviously reproduces TEC much better than IRI2007. On the other hand, IRI2007 presents better reproductions between $3.5\text{--}5.0 \times 10^{10}$ photon/cm² s, which is the moderate

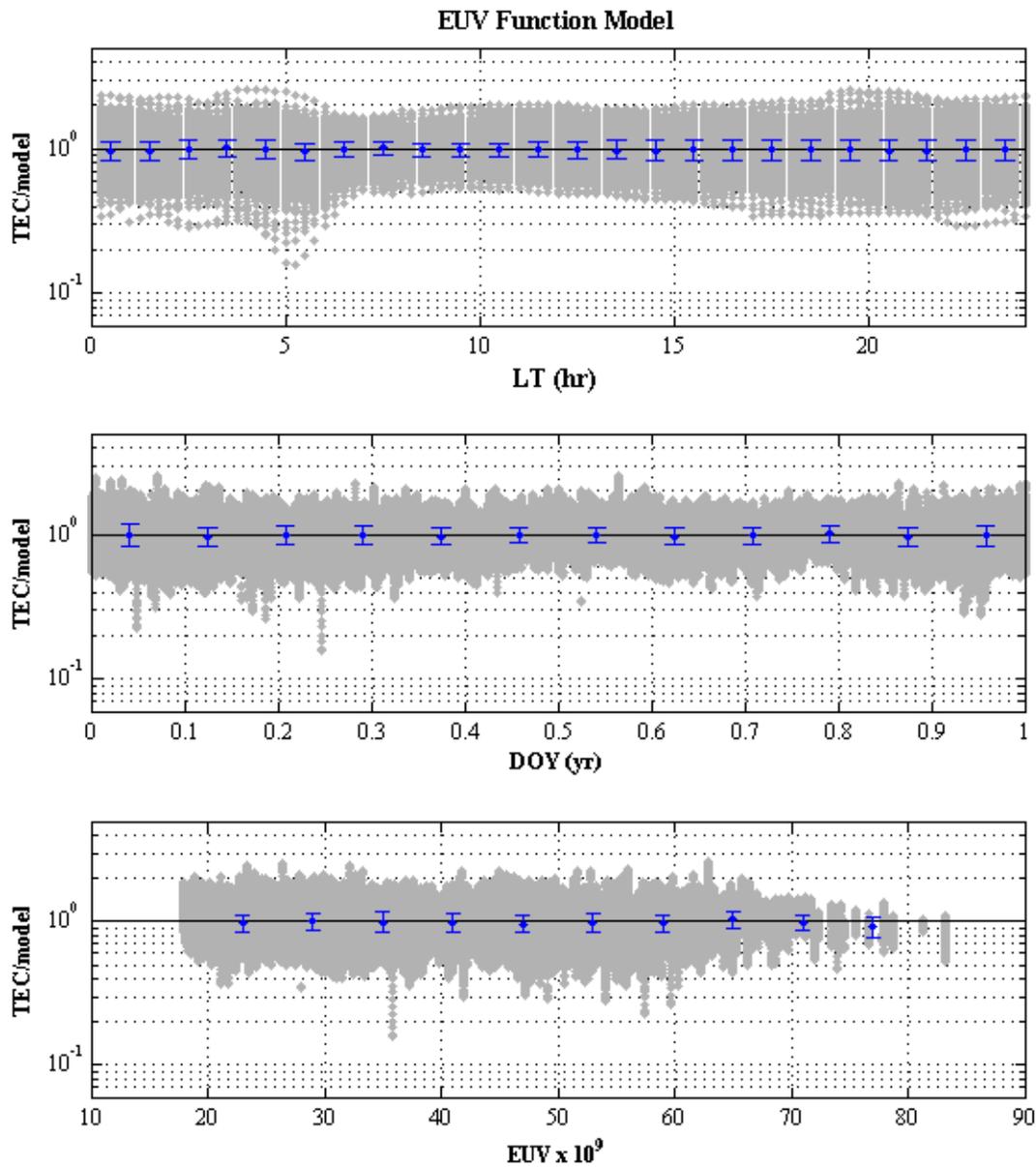


Fig. 9. Ratios of observed TEC to the F_{EUV} functional model (gray dots) and their medians (blue dots) with error bars during the magnetic quiet condition ($D_{st} > -30$ nT).

solar activity. This is the reason why RMSs of IRI2007 is comparable to the F_{EUV} function model and the 15-days running median during 1998–1999 and 2003–2004 (Fig. 8).

Comparison of RMS also indicates primacy of the F_{EUV} function model over IRI2007 (Fig. 11). RMSs of the F_{EUV} function model are always less than IRI2007 in any LT, DOY and EUV. Differences of RMS become large during daytime, equinoxes. Especially, the differences of RMS become large when the solar activity is high and low.

7 Concluding remarks

We constructed the empirical models of TEC based on median values and the functional fitting over the Taiwan area (120° E, 24° N) during geomagnetically quiet condition. The models provide TEC as functions of day of local time (LT), year (DOY) and several solar fluxes (F). This paper is for the first time comparing the median based model with the functional fitting model. Results indicate the functional fitting model yielding a better performance than the median based one. Meanwhile we find that the EUV radiation is essential to derive an optimal TEC. The F_{EUV} functional model can

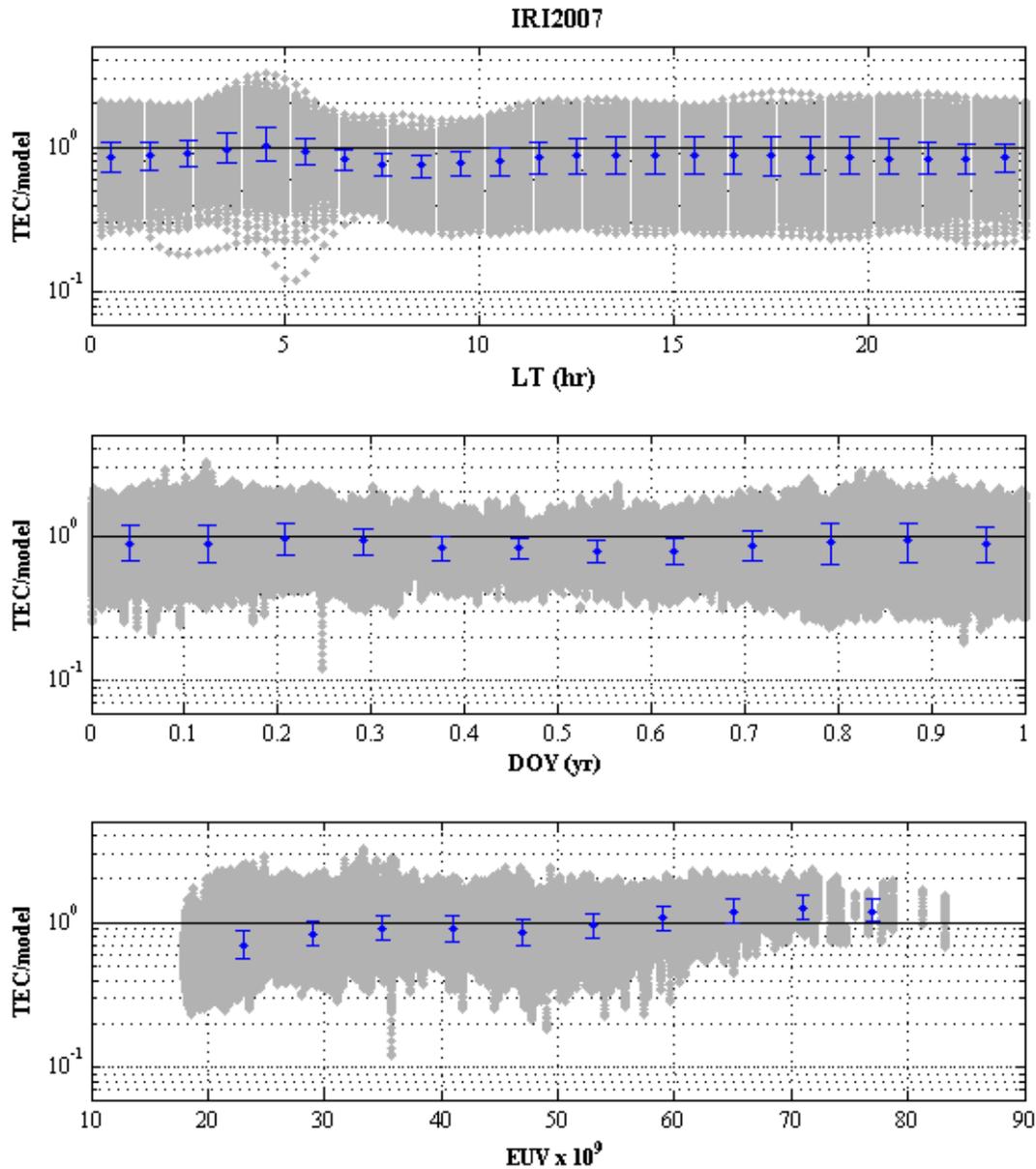


Fig. 10. Same as Fig. 10, but for IRI2007.

produce TEC the most accurately. Many authors have applied F_{81dc} as the indicator for EUV. However, our results suggest that F_{54dp} is better than F_{81dc} as the indicator not only for EUV but also for TEC over the Taiwan area. F_{EUV} is the best parameter under the quiet condition, but we must note that F_{EUV} cannot be used for the indicator of TEC when EUV is highly fluctuating due to the intensive flares. The 15-days running median is accurate as well as the F_{EUV} function model during the low solar activity. However, the 15-days running median get worse during the high solar activity as compared with the F_{EUV} functional model.

Though IRI2007 has been improved in the upper ionosphere from IRI2000, IRI2007 has still shortcoming in LT, DOY, and the solar activity variations. Especially, IRI2007

overestimates when the solar activity is low, and underestimates when the solar activity is high.

Our models can only reproduce TEC during the quiet magnetic condition. Comparing our models to observations during disturbed conditions, we can investigate effects which produced by the magnetic storms and the other phenomena, such as the earthquakes. Our model will be improved to add a storm time option in near future. Although the models have been constructed to achieve to detect the ionospheric perturbations associated with the large earthquakes, because regional empirical model in the place where the equatorial ionospheric anomaly is observed has not been constructed yet, our models must contribute to improve IRI.

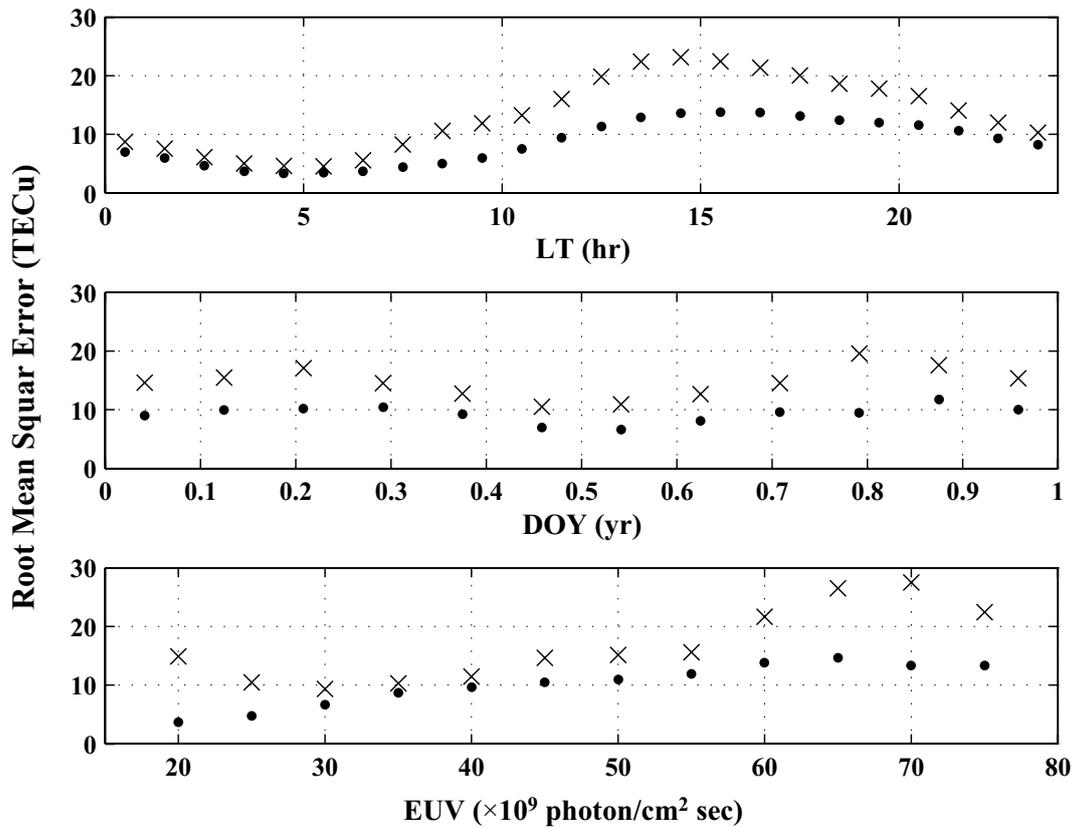


Fig. 11. Comparison of F_{EUV} functional model with IRI2007. Dots and crosses denote root mean square errors of F_{EUV} functional model and IRI2007, respectively.

Appendix A

Applied functions

Applied functions f' , g' , and h' are defined as follows:

$$f'_1 = 1, \tag{A1}$$

$$f'_2 = F, \tag{A2}$$

$$f'_3 = F^2, \tag{A3}$$

$$f'_4 = F^3, \tag{A4}$$

$$g'_1 = 1, \tag{A5}$$

$$g'_2 = \cos(2\pi \cdot \text{DOY}), \tag{A6}$$

$$g'_3 = \sin(2\pi \cdot \text{DOY}), \tag{A7}$$

$$g'_4 = \cos(4\pi \cdot \text{DOY}), \tag{A8}$$

$$g'_5 = \sin(4\pi \cdot \text{DOY}), \tag{A9}$$

$$g'_6 = \cos(6\pi \cdot \text{DOY}), \tag{A10}$$

$$g'_7 = \sin(6\pi \cdot \text{DOY}), \tag{A11}$$

$$h'_1 = 1, \tag{A12}$$

$$h'_2 = \cos(\pi \cdot \text{LT}/12), \tag{A13}$$

$$h'_3 = \sin(\pi \cdot \text{LT}/12), \tag{A14}$$

$$h'_4 = \cos(\pi \cdot \text{LT}/6), \tag{A15}$$

$$h'_5 = \sin(\pi \cdot \text{LT}/6), \tag{A16}$$

$$h'_6 = \cos(\pi \cdot \text{LT}/4), \tag{A17}$$

$$h'_7 = \sin(\pi \cdot \text{LT}/4), \tag{A18}$$

$$h'_8 = \cos(\pi \cdot \text{LT}/3), \tag{A19}$$

$$h'_9 = \sin(\pi \cdot \text{LT}/3), \tag{A20}$$

Appendix B

The coefficients α

The coefficients α for the F_{EUV} functional models. Labels out of the table show number of the coefficients α . Input parameter “EUV (0.1–50 nm)” must be multiplied by 10^9 when the F_{EUV} functional model is used. Please contact the corresponding author if you need electric data.

Table B1.

	1	2	3	4	5	6	7	8	9	10
0	4.57E+00	8.32E+00	-7.08E+00	1.45E+00	-7.97E+00	-2.07E+00	-2.51E-01	-2.07E-01	-6.96E-02	9.28E+00
10	-3.87E+00	-1.05E+01	-2.55E+00	1.64E+00	-3.03E+00	6.92E-01	-1.63E+00	3.22E-01	4.23E+00	-8.35E-01
20	-2.19E-01	5.46E+00	4.97E+00	1.31E+00	-3.56E+00	2.93E-01	3.33E-01	-9.63E+00	-1.17E+01	-1.16E+00
30	-5.26E+00	7.87E+00	1.82E+00	1.18E+00	8.15E-01	8.54E-01	-6.36E+00	-7.63E+00	-2.81E+00	6.57E+00
40	6.65E+00	-3.72E+00	-2.17E+00	6.73E-01	-8.91E-03	2.53E+00	1.88E+01	-2.05E+00	-6.03E+00	-5.59E+00
50	-3.27E-01	2.05E-02	1.92E+00	1.75E+00	4.09E-02	-7.10E+00	-2.97E+00	-3.06E+00	5.50E+00	1.69E-01
60	-3.41E-01	1.55E+00	3.21E+00	-7.16E-03	-9.65E-01	7.50E-01	-2.84E-02	7.82E-01	2.13E-01	-4.96E-02
70	5.47E-03	1.13E-02	-1.04E+00	6.06E-01	1.14E+00	4.36E-01	-1.48E-01	2.35E-01	-1.39E-01	1.44E-01
80	-3.56E-02	-4.18E-01	4.00E-02	1.18E-01	-3.95E-01	-4.58E-01	-1.58E-01	3.10E-01	-5.66E-03	-2.50E-02
90	9.11E-01	1.14E+00	-8.43E-02	3.87E-01	-8.32E-01	-2.15E-01	-5.83E-02	-5.49E-02	-6.22E-02	5.12E-01
100	6.61E-01	1.82E-01	-6.34E-01	-5.25E-01	3.50E-01	2.15E-01	-5.42E-02	-6.33E-03	-1.95E-01	-1.62E+00
110	1.22E-01	4.18E-01	4.63E-01	-6.11E-04	3.40E-02	-1.47E-01	-1.31E-01	4.61E-02	6.11E-01	1.54E-01
120	2.65E-01	-4.49E-01	-1.08E-02	2.38E-02	-1.48E-01	-2.95E-01	3.03E-02	8.08E-03	-3.83E-02	1.56E-03
130	-1.55E-02	-5.37E-03	2.80E-03	5.88E-04	8.61E-04	2.48E-02	-2.56E-02	-3.02E-02	-1.30E-02	4.71E-03
140	-6.84E-03	4.57E-03	-4.87E-03	1.62E-03	1.65E-02	-1.69E-03	-1.05E-02	8.32E-03	1.39E-02	5.22E-03
150	-8.03E-03	-2.06E-04	6.16E-04	-3.38E-02	-2.48E-02	1.16E-02	-1.90E-02	2.24E-02	7.45E-03	7.76E-04
160	4.77E-04	4.19E-04	-1.52E-02	-1.60E-02	-2.16E-03	1.78E-02	1.14E-02	-1.03E-02	-5.52E-03	1.52E-03
170	-2.32E-04	5.06E-03	4.35E-02	-2.79E-03	-9.47E-03	-1.20E-02	1.19E-03	-1.26E-03	3.69E-03	2.57E-03
180	-4.12E-03	-1.59E-02	-6.81E-05	-7.30E-03	1.11E-02	-7.69E-05	-6.57E-04	4.36E-03	8.17E-03	-2.34E-04
190	-9.01E-07	3.16E-04	-2.92E-06	7.82E-05	4.01E-05	-3.24E-05	-1.51E-05	-8.72E-06	-1.49E-04	2.32E-04
200	1.91E-04	1.04E-04	-4.63E-05	5.27E-05	-4.13E-05	4.20E-05	-1.52E-05	-1.74E-04	3.04E-05	1.52E-04
210	-4.35E-05	-1.32E-04	-5.01E-05	6.30E-05	4.70E-06	-7.32E-06	2.83E-04	1.47E-04	-1.17E-04	6.63E-05
220	-1.70E-04	-6.90E-05	6.19E-06	4.81E-06	2.28E-06	1.25E-04	1.29E-04	-4.69E-06	-1.54E-04	-7.02E-05
230	9.07E-05	4.16E-05	-1.50E-05	5.58E-06	-5.32E-05	-3.60E-04	3.62E-05	7.29E-05	1.05E-04	-1.40E-05
240	1.21E-05	-2.70E-05	-1.53E-05	6.52E-05	1.23E-04	-3.43E-05	6.15E-05	-8.46E-05	2.57E-06	4.16E-06
250	-3.89E-05	-7.10E-05								

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