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Abundances of Recently Discovered Planetary Nebulae Towards the Galactic Bulge^{*}

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Abstract. In this work we report spectrophotometric observations of a planetary nebula sample towards the galactic bulge. A total of 45 PNe was observed and their physical parameters (electron density and temperature) were derived. Ionic abundances were calculated using a three-level atom model and abundances were derived through ionisation correction factors. Results show low abundance objects at high galactic latitudes, indicating a possible vertical gradient inside the bulge. A few objects with low N/O ratio were found, which could have originated from old, low mass progenitors.

Key words: planetary nebulae – galactic bulge – spectroscopy

1. Introduction

The galactic bulge is a structure whose properties have been extensively studied in recent years. In particular, abundance studies show many aspects of bulge populations, sometimes contradictory: while earlier works (e.g. Rich 1988) indicate abundances up to $[\text{Fe}/\text{H}] \sim 0.30$, other more recent studies (McWilliam & Rich 1994, Ibata & Gilmore 1995) indicate subsolar values, derived from K giants. Using spectra of integrated light from the galactic bulge in Baade's window, Idiart et al. (1996) derive a mean metallicity $[\text{Fe}/\text{H}] = -0.02$, which is consistent with the mean values found by Rich (1988) and McWilliam & Rich (1994). Comprehensive revisions of abundance data and the current ideas on Galaxy evolution can be found in van den Bergh (1996) and McWilliam (1997).

In this context, abundances from planetary nebulae (PNe) play an important role. The derivation of nebular abundances of light elements like O, S, Ar from the flux of forbidden emission lines is quite convenient at the large distances and strong reddening of the galactic bulge.

Abundances from these objects reflect the state of the interstellar medium at the progenitor formation epoch, and therefore can be related directly to stellar abundances. On the other hand, He and N abundances are modified along the stellar evolution of the progenitor and can be related to their masses (and age).

Some previous works (Ratag et al. 1992, 1997; Cuisinier et al. 2000) give abundances of different samples of bulge PNe. In both cases, objects were selected from the Acker et al. (1992) catalogue, and results indicate an abundance pattern comparable to their disk counterparts. A point yet to be tackled is the possibility of the existence of a chemical abundance gradient within the bulge. There are several suggestions for the existence of a chemical abundance gradient within the bulge, as proposed by Minitti et al. (1995), Frogel et al. (1999) or Tiede et al. (1995). In this last work they found evidence of a gradient in Fe/H for the region $-12 \leq b \leq -3$.

As part of a long-term program of derivation of nebular abundances, we have carried out observations of southern planetary nebulae (see, e.g. Costa et al. 1996, 2000 and references therein) aiming to derive their nebular physical parameters and chemical abundances. In this work, we report spectroscopic observations for a sample of 44 PNe towards the galactic bulge. From this sample, chemical abundances were derived. Objects were selected from new samples recently published (Kohoutek 1994; Beaulieu et al. 1999), and with these results we intend to enlarge the sample of bulge PNe with known chemical abundances, to obtain the abundance distribution of the intermediate mass star population in the galactic bulge. These results give important clues for the chemical enrichment and evolution of the galactic bulge.

In section 2 we present our observation and reduction procedures, in sect. 3 physical parameters and chemical abundances derivation are derived, in sect. 4 results are discussed and general conclusions are drawn.

2. Observations and data reduction

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^{*} Based on observations made at the European Southern Observatory (Chile) and Laboratório Nacional de Astrofísica (Brasil)

2.1. Observations

Observations were made at the 1.60m telescope of the National Laboratory for Astrophysics (LNA, Brasópolis, Brazil) and at the 1.52m of the European Southern Observatory (ESO/La Silla, Chile). In both telescopes, Cassegrain Boller & Chivens spectrographs were used, with reciprocal dispersions of 4.4 Å/pixel at LNA and 2.2 Å/pixel at ESO. A long east-west slit was used in all observations. In both observatories, slits of 2 arcsecs width were used. The log of the observations and identification of the objects are given in table 1. Each object was observed at least two times and line fluxes were derived from the average of all measures for the same object.

Data reduction was performed using the IRAF package, following the standard procedure for longslit spectra: correction of bias, flat-field, extraction, wavelength and flux calibration. For flux calibration, we observed at least three spectrophotometric standard stars each night. Atmospheric extinction was corrected through mean coefficients derived for each observatory.

Line fluxes were calculated adopting gaussian profiles, and a gaussian de-blending routine was used when necessary. Line fluxes are given in tables 2 and 3 using the usual scale of $F(H\beta)=100$. The $H\beta$ fluxes through the slit, in physical units ($\text{erg}/\text{cm}^2.\text{s}$), and corrected for interstellar and atmospheric reddening are listed in table 4. Interstellar reddening was derived from the Balmer ratio $H\alpha/H\beta$, assuming Case B (Osterbrock 1989), using the interstellar extinction law by Cardelli et al. (1989) and adopting $R_V=3.1$ for all objects. $E(B-V)$ values derived for the nebulae are given in table 4.

All objects from Beaulieu et al. (1999) and Kohoutek's (1994) table 5 were observed. Those not present in table 1 do not have a S/N satisfactory for plasma diagnostics. In particular, objects K5-8, SB-05, SB-27 and SB-29 do not seem to be PNe.

2.2. Physical parameters

Electron densities were derived from the [SII] $\lambda 6716/\lambda 6731$ ratio, while electron temperatures were derived from both [OIII] $\lambda 4363/\lambda 5007$ and [NII] $\lambda 5755/\lambda 6584$ line ratios. When both temperatures were available, the [OIII] temperature was used to derive ionic abundances of species with higher ionisation potentials as O^{+2} , Ar^{+2} , Ar^{+3} , Ne^{+2} , and the [NII] temperature was used for lower potential lines such as O^+ , N^+ , S^+ . When only one temperature was available, it was used for all ions. Concerning densities, when [SII] lines could not be used for diagnostics, an average value of 5000 cm^{-3} was adopted. Derived values are shown in table 4. Electron densities are in units of 10^3 cm^{-3} and temperatures in units of 10^4 K .

One of the main problems of observations towards the bulge is the high extinction, which affects mainly lower in-

Table 1. Log of the observations

PNe	PN G / PK	Num.	Date of Obs.	Place
K5-1	PK 000+04.3	2	20-Jul-99	LNA
K5-3	PN G 002.6+05.5	2	31-May-98	LNA
		2	16-Aug-99	ESO
K5-4	PN G 351.9-01.9	2	31-May-98	LNA
		2	15-Aug-99	ESO
K5-5	PN G 001.5+03.6	2	20-Jul-99	LNA
K5-6	PN G 003.6+04.9	2	20-Jul-99	LNA
K5-7	PN G 003.1+04.1	3	19-Jul-99	LNA
		2	16-Aug-99	ESO
K5-9	PK 355-01.1	2	19-Jul-99	LNA
K5-11	PK 002+02.1	2	31-May-98	LNA
		2	15-Aug-99	ESO
K5-12	PK 353-03.1	2	19-Jul-99	LNA
K5-13	PK 002+02.2	2	31-May-98	LNA
		2	16-Aug-99	ESO
K5-14	PN G 003.9+2.6	1	18-Jul-99	LNA
		1	19-Jul-99	LNA
K5-16	PN G 002.8+01.8	1	31-May-98	LNA
		2	18-Jul-99	LNA
K5-17	PN G 004+02.2	2	18-Jul-99	LNA
K5-19	PN G 005.1+02.0	1	18-Jul-99	LNA
K5-20	PN G 356.8-03.0	3	18-Jul-99	LNA
		2	16-Aug-99	ESO
SB 01	PN G 000.1-08.0	2	24-Jun-00	LNA
SB 02	PN G 000.5-05.3	2	21-Jun-00	LNA
SB 03	PN G 000.7-06.1	2	23-Jun-00	LNA
SB 04	PN G 001.1-06.4	2	24-Jun-00	LNA
SB 06	PN G 001.6-05.9	2	23-Jun-00	LNA
SB 12	PN G 005.4-06.1	2	24-Jun-00	LNA
SB 15	PN G 009.3-06.5	2	23-Jun-00	LNA
SB 17	PN G 011.1-07.9	2	17-Aug-99	ESO
SB 18	PN G 011.4-07.3	1	23-Jun-00	LNA
		2	25-Jun-00	LNA
SB 19	PN G 014.4-06.1	2	17-Aug-99	ESO
SB 20	PN G 014.8-08.4	2	22-Jun-00	LNA
SB 21	PN G 016.0-07.6	2	16-Aug-99	ESO
SB 24	PN G 017.5-09.2	2	21-Jun-00	LNA
SB 25	PN G 341.0+09.4	2	19-Aug-99	ESO
SB 26	PN G 341.7-06.0	2	21-jun-00	LNA
SB 28	PN G 342.3-06.0	2	18-Aug-99	ESO
SB 30	PN G 343.9-05.8	2	23-Jun-00	LNA
SB 31	PNG 347.9-06.0	2	18-Aug-99	ESO
SB 32	PNG 349.7-09.1	2	20-Aug-99	ESO
		1	21-Jun-00	LNA
SB 33	PNG 351.2-06.3	2	24-Jun-00	LNA
SB 34	PNG 351.5-06.5	1	18-Aug-99	ESO
		1	19-Aug-99	ESO
SB 35	PNG 351.7-06.6	2	20-Aug-99	ESO
		2	24-Jun-00	LNA
SB 37	PNG 352.6-04.9	2	19-Aug-99	ESO
SB 38	PNG 352.7-08.4	2	22-Jun-00	LNA
SB 42	PNG 355.3-07.5	2	22-Jun-00	LNA
SB 44	PNG 356.0-07.4A	2	18-Aug-99	ESO
SB 50	PNG 357.3-06.5	2	22-Jun-00	LNA
SB 52	PNG 358.3-07.3	2	22-Jun-00	LNA
SB 53	PNG 358.7-05.1	2	23-Jun-00	LNA
SB 55	PNG 359.4-08.5	2	17-Aug-99	ESO

Table 2. Line fluxes

Line	K5-1	K5-3	K5-4	K5-5	K5-6	K5-7	K5-9	K5-11	K5-12	K5-13	K5-14	K5-16
[OII] λ 3726+29	-	63.0	133.7	-	-	-	-	-	-	117.3	-	-
[NeIII] λ 3869	-	126.1	172.6	-	-	118.7	-	-	51.9	196.3	38.1	-
[NeIII] λ 3967	-	55.3	62.1	-	-	-	-	-	23.4	75.8	-	-
$H\gamma$ λ 4340	65.5	56.2	57.3	91.6	65.5	74.6	-	-	75.5	54.2	41.7	-
[OIII] λ 4363	7.9	12.9	6.9	30.6	25.8	26.8	-	2.6	9.8	20.3	37.5	-
HeI λ 4472	11.9	10.1	6.2	18.6	10.2	-	8.7	-	-	11.8	4.2	-
HeII λ 4686	-	51.5	-	-	102.2	106.6	18.3	21.4	71.9	73.8	26.2	15.8
$H\beta$ λ 4861	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0
[OIII] λ 4959	168.7	276.7	357.9	283.0	310.6	339.6	225.4	172.1	402.8	441.6	571.2	287.1
[OIII] λ 5007	507.7	818.9	1061.5	843.3	915.5	1024.7	659.3	496.4	1201.3	1328.9	1663.8	859.1
HeII λ 5412	-	4.6	0.4	-	13.6	6.0	1.9	-	5.4	6.6	2.4	2.3
[NII] λ 5755	0.9	0.5	1.9	2.0	2.9	-	0.7	3.1	0.7	4.3	6.4	6.1
HeI λ 5876	17.5	11.6	17.4	9.7	3.6	8.6	21.1	25.1	7.4	11.8	15.5	18.3
[OI] λ 6300	0.6	1.0	6.2	2.9	1.7	12.9	2.3	8.3	2.4	6.4	14.0	31.1
[SIII] λ 6312	0.7	1.4	1.8	1.3	1.8	-	0.6	1.1	1.9	6.0	3.9	5.0
[NII] λ 6548	4.2	10.7	25.4	3.7	2.3	9.5	11.7	107.9	6.8	111.5	58.1	164.8
$H\alpha$ λ 6563	285.2	285.2	285.2	285.3	285.2	285.2	285.3	285.2	285.2	285.2	285.2	285.3
[NII] λ 6584	18.1	23.8	71.8	20.5	7.4	19.1	41.8	333.5	25.1	345.7	208.6	501.8
HeI λ 6678	5.1	3.7	4.4	4.2	2.4	4.6	6.4	7.6	2.4	5.2	4.4	5.7
[SII] λ 6716	1.4	3.4	3.2	1.5	2.5	5.4	2.7	25.4	2.8	28.9	8.6	59.5
[SII] λ 6731	1.9	4.4	6.1	3.0	2.9	4.6	4.4	26.8	5.0	35.9	16.6	62.3
[ArIII] λ 7135	15.5	11.0	15.3	10.4	8.9	24.7	16.1	19.5	91.4	40.4	32.3	24.5
[OII] λ 7320+30	2.4	2.8	11.7	10.0	2.0	3.7	3.3	6.0	3.3	5.2	26.5	10.8
Line	K5-17	K5-19	K5-20	SB01	SB02	SB03	SB04	SB06	SB12	SB15	SB17	SB18
[OII] λ 3726+29	-	-	33.1	109.0	81.7	41.2	-	458.8	-	56.8	553.6	154.3
[NeIII] λ 3869	-	-	87.3	152.2	25.9	89.1	78.0	82.2	-	189.6	-	70.7
[NeIII] λ 3967	-	-	48.4	81.3	50.8	-	-	75.4	-	-	19.5	35.7
$H\gamma$ λ 4340	32.4	-	47.3	35.8	53.4	49.1	49.3	36.5	37.3	46.5	53.8	35.5
[OIII] λ 4363	23.7	54.4	7.0	7.6	8.6	36.5	15.4	19.0	35.9	46.7	-	27.6
HeI λ 4472	-	-	7.9	22.4	18.1	3.8	-	12.5	-	-	-	12.7
HeII λ 4686	-	95.8	-	37.7	32.2	59.5	21.6	50.4	72.1	51.3	-	64.3
$H\beta$ λ 4861	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0
[OIII] λ 4959	408.2	422.1	269.7	164.9	159.9	208.6	181.0	154.8	322.6	442.5	-	339.2
[OIII] λ 5007	1247.1	1173.4	808.3	540.1	477.8	616.5	522.1	504.5	799.8	1300.1	13.4	972.8
HeII λ 5412	0.9	-	-	-	-	3.7	-	6.6	-	27.0	-	-
[NII] λ 5755	-	1.6	1.8	-	2.6	1.6	-	-	-	-	6.5	-
HeI λ 5876	14.1	10.4	17.6	6.7	18.3	14.2	11.7	10.1	11.9	11.8	18.6	6.0
[OI] λ 6300	2.4	-	3.1	7.0	13.6	-	-	-	-	-	37.4	3.7
[SIII] λ 6312	1.2	-	0.5	-	-	-	-	-	-	-	-	-
[NII] λ 6548	5.7	13.5	2.0	6.7	28.4	12.1	42.7	176.5	-	-	94.5	-
$H\alpha$ λ 6563	285.2	285.2	285.2	285.1	285.1	285.1	285.0	285.1	285.0	285.1	285.1	285.0
[NII] λ 6584	20.8	42.8	4.5	17.5	78.5	42.4	116.4	531.6	7.5	1.6	256.6	3.6
HeI λ 6678	3.7	2.5	4.6	5.0	4.6	5.6	4.8	6.6	17.2	8.8	6.0	2.8
[SII] λ 6716	2.0	7.2	0.8	5.0	12.7	7.0	15.1	14.4	-	-	57.5	1.8
[SII] λ 6731	4.0	8.5	0.8	6.6	10.3	6.2	11.7	10.9	-	6.0	44.1	1.5
[ArIII] λ 7135	11.1	14.9	8.2	4.3	10.9	10.5	8.2	20.8	15.2	2.2	1.2	2.9
[OII] λ 7320+30	4.8	3.2	2.9	4.0	5.8	-	5.2	18.1	-	8.9	13.0	5.6

tensity lines. Lines [SII] λ 6716+31 for example, are sometimes weak, but needed to derive electron densities. Other weak lines like [SIII] λ 6312 or HeII λ 5412 are also needed to derive the appropriate ionic abundances. A more serious problem arises from [OIII] λ 4363 and [NII] λ 5755 because

abundances are strongly sensitive to electron temperature. Errors in these line fluxes will contribute to the derived errors in the final abundances.

Table 3. Line fluxes (continued)

Line	SB19	SB20	SB21	SB24	SB25	SB26	SB28	SB30	SB31	SB32	SB33	SB34
[OII] λ 3726+29	42.4	59.8	-	39.6	209.1	29.2	-	53.1	14.4	398.6	103.7	-
[NeIII] λ 3869	63.1	254.1	51.9	-	97.1	126.3	55.9	74.8	84.5	60.2	24.5	79.6
[NeIII] λ 3967	-	-	30.6	35.8	-	53.9	25.5	32.3	36.2	34.2	-	-
H γ λ 4340	31.9	36.4	54.6	21.6	49.9	40.1	66.2	52.1	63.6	47.0	51.1	68.5
[OIII] λ 4363	2.9	70.5	14.3	6.9	7.0	21.0	29.3	16.1	14.3	15.9	22.9	22.1
HeI λ 4472	7.6	-	-	-	-	-	-	3.6	-	9.4	28.9	-
HeII λ 4686	-	30.2	116.9	90.7	-	94.5	116.2	46.6	117.0	25.7	17.5	140.4
H β λ 4861	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0
[OIII] λ 4959	224.7	497.0	188.6	117.5	170.5	308.7	174.5	320.5	227.2	189.6	222.7	162.9
[OIII] λ 5007	648.7	1441.1	571.9	384.9	515.8	915.6	525.0	941.1	662.4	565.2	585.0	509.8
HeII λ 5412	-	-	7.9	3.2	-	7.4	6.7	3.6	7.3	-	-	21.5
[NII] λ 5755	-	-	-	-	-	-	-	-	-	5.0	11.9	-
HeI λ 5876	22.5	13.7	4.0	5.8	18.1	3.0	1.1	10.5	1.3	15.1	10.7	2.4
[OI] λ 6300	-	-	-	-	9.3	2.4	-	4.2	0.9	32.1	31.8	7.9
[SIII] λ 6312	1.5	-	-	-	1.0	0.4	-	0.6	1.8	8.6	-	-
[NII] λ 6548	-	-	-	-	107.0	2.5	-	6.4	-	37.6	77.0	7.9
H α λ 6563	285.1	285.0	285.1	285.1	285.1	285.1	285.1	285.1	285.1	285.0	285.0	285.1
[NII] λ 6584	8.2	1.8	1.6	6.3	313.6	1.7	1.3	21.6	4.8	110.6	205.2	6.4
HeI λ 6678	8.5	2.8	1.4	6.8	4.2	2.5	-	2.8	1.8	7.9	10.2	1.3
[SII] λ 6716	2.7	3.9	-	2.1	38.4	-	-	1.2	1.0	32.8	32.3	-
[SII] λ 6731	3.7	2.6	-	4.0	31.9	-	-	2.2	1.2	26.0	28.9	-
[ArIII] λ 7135	19.8	3.0	7.8	4.8	19.6	5.4	4.4	5.3	5.5	8.8	15.7	4.6
[OII] λ 7320+30	14.6	-	3.0	10.3	-	1.0	-	6.0	1.2	15.9	12.9	-

Line	SB35	SB37	SB38	SB42	SB44	SB50	SB52	SB53	SB55
[OII] λ 3726+29	29.8	38.4	6.0	35.8	-	72.5	33.4	43.0	15.4
[NeIII] λ 3869	73.0	132.2	63.1	29.1	25.1	62.9	68.2	71.1	56.6
[NeIII] λ 3967	18.9	47.3	13.4	26.9	-	-	37.5	39.0	24.0
H γ λ 4340	42.6	76.2	44.6	32.8	32.4	51.3	56.1	53.2	46.1
[OIII] λ 4363	8.3	32.5	26.1	3.0	6.7	20.0	12.1	23.7	17.0
HeI λ 4472	11.9	-	-	7.6	-	-	2.8	-	-
HeII λ 4686	-	166.3	65.6	4.9	129.8	96.5	121.4	128.1	126.7
H β λ 4861	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0	100.0
[OIII] λ 4959	266.2	401.8	303.1	89.0	188.0	228.5	228.6	312.7	123.7
[OIII] λ 5007	766.0	1213.7	888.2	245.0	627.5	685.7	674.9	904.5	353.0
HeII λ 5412	-	12.0	3.5	-	7.6	5.9	9.8	8.7	9.0
[NII] λ 5755	1.1	-	-	0.3	-	-	0.2	1.2	-
HeI λ 5876	16.9	-	15.6	6.5	4.0	7.4	4.8	2.5	7.2
[OI] λ 6300	-	0.7	-	-	-	-	1.3	3.3	-
[SIII] λ 6312	-	3.4	-	-	-	-	1.2	4.2	-
[NII] λ 6548	3.4	-	-	-	15.1	-	0.8	5.6	-
H α λ 6563	285.1	285.1	285.0	285.1	285.0	285.0	285.0	285.1	285.1
[NII] λ 6584	2.7	13.9	2.2	10.8	19.8	14.0	5.6	16.1	3.2
HeI λ 6678	5.5	3.3	3.3	4.6	2.6	5.1	2.8	2.6	3.8
[SII] λ 6716	1.5	3.5	3.1	2.0	3.8	6.5	2.2	4.1	0.3
[SII] λ 6731	-	3.2	-	1.5	4.3	5.8	2.6	4.2	0.4
[ArIII] λ 7135	7.8	-	6.1	3.4	14.3	12.1	12.2	19.1	5.6
[OII] λ 7320+30	2.2	0.8	3.5	2.8	7.9	-	2.9	-	5.4

Examining the line fluxes for objects measured in both observatories, no systematic differences were found, and we plot all the data together .

As we have independent measures for each object, it was possible to estimate the errors in line intensities. For

this, we plot in fig. 1 all line fluxes for all objects using reddened values, and estimate the error for each line from its average and dispersion. From this distribution, we derived the following expression for the correlation between errors and fluxes:

$$\log(\Delta F) = -5.67 \pm 0.55 + (0.67 \pm 0.04)\log(F)$$

Where line fluxes are represented by F and their errors by ΔF (in physical units). The full line in fig. 1 represents $\Delta F = F$.

In order to estimate the errors in abundances, each line flux had its value varied randomly 500 times within its respective error interval, as well as the colour excesses, within the interval given in table 4. Physical parameters and abundances were then derived for each one. A histogram was built with these values and the final adopted error for each abundance was the full width of the gaussian profile fitted to each histogram, taken at half maximum. These errors are listed in table 6, together with the elemental abundances.

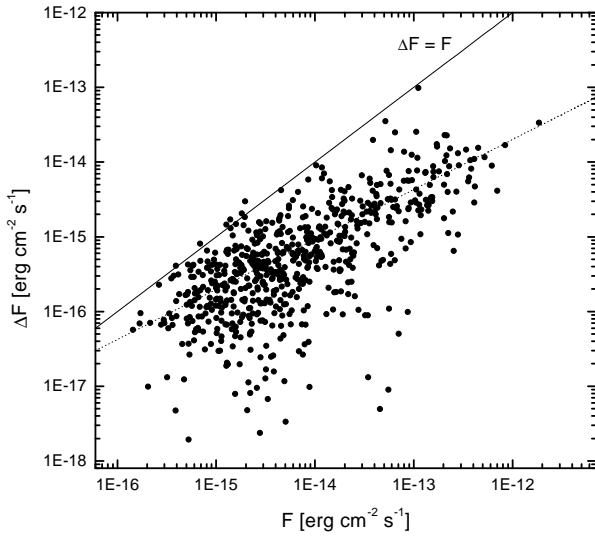


Fig. 1. Errors in line fluxes with respect to their values (in physical units).

This process was adopted to derive errors in abundances and temperatures. For densities the value distribution is not gaussian and errors were derived simply by averaging those derived from different measures for each object. Errors in reddening correspond to the dispersion of those calculated from each measure.

We noted from our dereddened fluxes that for some objects the $H\gamma$ flux is considerably different from the recombination value. This effect is stronger for low latitude objects, indicating a probable relation with variations in the total-to-selective absorption ratio R_V for some light paths towards the bulge. Stasińska et al. (1992) have already mentioned the possibility of different R_V values for bulge planetary nebulae. Variations in R_V , as pointed out

by Cardelli et al. (1989) can have considerable effects in the blue-UV region. As the $[OIII]\lambda 4363$ line is close to $H\gamma$ and important to determine electron temperatures, it was necessary to check how this effects our derived physical parameters, and therefore the chemical abundances. To verify, we scaled the $[OIII]\lambda 4363$ line using the ratio between recombination and measured value for $H\gamma$ as a correction factor, and rederived electron temperatures and chemical abundances for all objects. Figure 2 displays a comparison between original and rederived oxygen abundances. As can be seen, variations are within our expected errors. Chemical abundances were then rederived using these "scaled" temperatures and the procedure described below, and the results were found to be consistent with the original values within the expected errors.

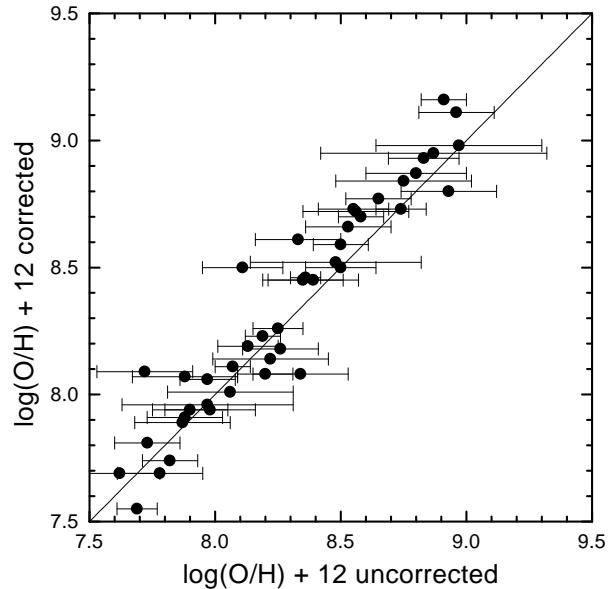


Fig. 2. Comparison between measured and rescaled oxygen abundances (see section 2.2).

3. Chemical abundances

Ionic abundances were calculated for the ions present in the optical spectra by solving the statistical equilibrium equations for a three-level atom model, including radiative and collisional transitions. Elemental abundances were then derived adopting ionisation correction factors (icf's) to account for unobserved ions of each element. Atomic data and icf's used in this work were the same as those adopted by Costa et al. (1996). Derived ionic abundances are listed in table 5, and the resulting elemental abundances are given in table 6.

Table 4. Physical parameters

PNe	E(B-V)	$\sigma_{E(B-V)}$	$\log(H\beta)$	n[SII]	$\sigma_{n[SII]}$	T[NII]	$\sigma_{T[NII]}$	T[OIII]	$\sigma_{T[OIII]}$
K 5-1	1.89	0.09	-11.22	1.4	0.2	1.82	0.56	1.37	0.25
K 5-3	1.27	0.13	-11.82	1.3	0.4	1.07	0.23	1.38	0.15
K 5-4	1.87	0.09	-10.84	7.6	1.8	1.17	0.16	0.99	0.10
K 5-5	2.24	0.11	-10.98	10.9	-	2.94	1.43	2.08	0.46
K 5-6	1.40	0.07	-11.98	1.0	0.3	-	-	1.82	0.25
K 5-7	1.43	0.14	-12.21	0.3	0.3	-	-	1.75	0.30
K 5-9	2.03	0.10	-11.23	3.5	1.5	1.01	0.24	-	-
K 5-11	1.92	0.38	-11.52	0.7	0.2	0.81	0.10	0.94	0.20
K 5-12	1.27	0.06	-11.76	5.5	1.6	1.23	0.24	1.07	0.09
K 5-13	1.82	0.18	-11.65	1.2	0.1	0.90	0.11	1.36	0.25
K 5-14	1.42	0.14	-11.67	8.7	-	1.24	0.17	1.60	0.17
K 5-16	1.98	0.40	-11.42	0.7	0.1	0.90	0.10	-	-
K 5-17	1.53	0.08	-11.42	11.6	-	-	-	1.47	0.15
K 5-19	1.85	0.19	-11.68	1.0	-	1.49	0.41	2.48	0.57
K 5-20	1.41	0.14	-11.83	0.5	-	-	-	1.10	0.15
SB 01	0.71	0.07	-12.99	1.4	1.3	-	-	1.33	0.35
SB 02	0.63	0.06	-12.63	0.2	0.1	1.41	0.22	1.46	0.17
SB 03	0.52	0.05	-12.67	0.3	-	1.50	0.26	3.09	0.38
SB 04	0.43	0.04	-13.26	0.1	0.1	-	-	1.86	0.27
SB 06	0.47	0.05	-13.30	0.1	-	-	-	2.20	0.37
SB 12	0.11	0.11	-14.58	-	-	-	-	2.35	0.61
SB 15	0.42	0.21	-13.07	-	-	-	-	2.08	0.23
SB 17	0.58	0.06	-13.25	0.1	-	1.22	0.19	-	-
SB 18	0.31	0.09	-13.49	0.3	-	-	-	1.82	0.22
SB 19	0.49	0.05	-13.18	1.8	-	-	-	0.90	0.11
SB 20	0.40	0.08	-13.27	-	-	-	-	2.58	0.29
SB 21	0.47	0.05	-12.96	-	-	-	-	1.70	0.20
SB 24	0.37	0.02	-13.68	9.1	-	-	-	1.45	0.25
SB 25	1.05	0.11	-12.86	0.2	0.1	-	-	1.30	0.29
SB 26	0.60	0.03	-12.70	-	-	-	-	1.62	0.16
SB 28	0.51	0.05	-13.44	-	-	-	-	2.90	0.64
SB 30	0.47	0.02	-12.05	5.5	2.8	-	-	1.42	0.07
SB 31	0.70	0.04	-12.40	1.0	0.5	-	-	1.58	0.14
SB 32	0.18	0.04	-13.70	0.2	0.2	1.72	0.40	1.82	0.25
SB 33	0.50	0.05	-13.56	0.4	0.1	2.06	0.59	2.18	0.38
SB 34	0.56	0.11	-13.23	-	-	-	-	2.39	0.44
SB 35	0.49	0.05	-12.76	-	-	-	-	1.18	0.10
SB 37	0.95	0.05	-11.91	0.4	0.1	-	-	1.78	0.13
SB 38	0.04	0.04	-13.65	-	-	-	-	1.85	0.19
SB 42	0.48	0.02	-12.35	0.1	0.1	1.31	0.28	1.24	0.11
SB 44	0.11	0.06	-13.73	0.9	0.3	-	-	1.19	0.15
SB 50	0.43	0.09	-13.12	0.3	0.3	-	-	1.86	0.21
SB 52	0.32	0.03	-12.76	1.0	1.2	1.46	0.48	1.46	0.10
SB 53	0.50	0.05	-12.63	0.7	0.1	2.53	0.74	1.74	0.13
SB 55	0.46	0.23	-12.95	1.7	-	-	-	2.56	0.50

Oxygen abundance requires ionic abundances of O^+ and O^{++} . While there are strong [OIII] lines at $\lambda 4959+5007 \text{ \AA}$, for O^+ we have in principle two possibilities: the blue pair [OII] $\lambda 3626+29 \text{ \AA}$, and the red one [OII] $\lambda 7320+20 \text{ \AA}$. In this work we adopted the red pair due to the greater sensitivity to reddening of the blue-UV region, as discussed in section 2.2, and due to the greater efficiency of the instrumental set in the red re-

gion. For many objects we were not able to get flux for the $\lambda 3626+29$ pair due to low S/N. In view of this fact we chose to keep the same line for all objects. Only for a few objects, as can be seen in tables 2 and 3, do we not have the red pair and were forced to adopt the blue one to derive O^+ abundance.

To verify an eventual dependence of the line adopted on the O^+ abundance, for those objects having both blue

and red lines measured we calculated the ionic abundance from both and compare them. The result is shown in figure 3. As can be seen, both values are perfectly compatible, and we expect no differences in the final oxygen abundance when using blue or red [OII] lines.

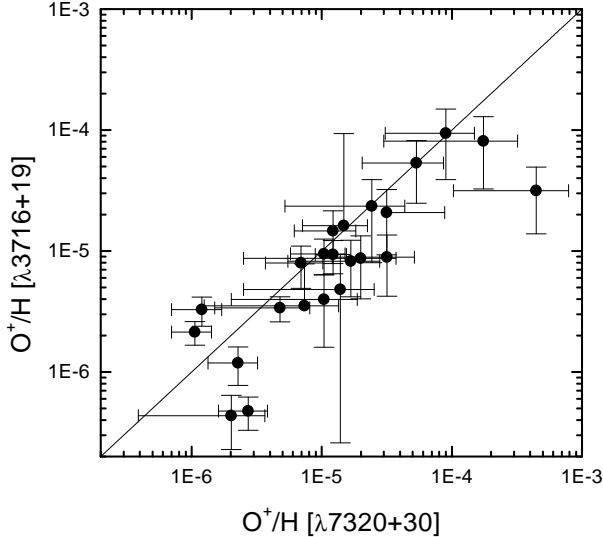


Fig. 3. O^+ ionic abundance derived from blue and red [OII] pairs (see text). Both values are clearly compatible, the diagonal $x=y$ line was drawn over the points and does not represent any adjustment.

Derivation of helium abundances requires special caution because collisional effects are present. In the literature there are correction factors to be applied to HeI abundances, however they may be overestimated up to a factor of two (Peimbert & Torres-Peimbert (1987). Here we adopted the recombination theory for HeI and HeII, using recombination coefficients from Péquignot (1991). HeI abundance was also corrected for collisional effects using the correction terms from Kingdon & Ferland (1995).

Helium abundance is given as He/H and other elemental abundances are given as $\varepsilon(X)=\log(X/H)+12$. Errors in the abundances refer always to X/H .

4. Discussions and Conclusion

Different samples were used for this work. Most of the objects listed by Kohoutek (1994) are in the northern galactic hemisphere while those from Beaulieu et al. (1999) are mostly in the southern galactic hemisphere. Also, we selected additional data from the literature for our analysis. These additional abundances are not homogeneous with respect to observation and reduction procedures, however they are important for the sampling of the range of elemental abundances. There are few works on bulge PNe

abundances, as this kind of result requires a good S/N ratio towards a region of high extinction. Three works were then selected: Cuisinier et al. (2000), Webster (1988) and Aller & Keyes (1987). Webster's abundances were re-derived using the same icf's as we used for ours. For each of these samples we selected only objects with good spectra.

Table 6 lists our final abundances for each object. Mean values and their distribution were also derived for the sample, where distribution is represented by a 2σ deviation from the mean value. Final values are: He: 0.126 ± 0.054 ; N: 7.64 ± 1.10 ; O: 8.22 ± 0.85 ; S: 6.48 ± 1.87 ; Ar: 5.95 ± 1.01 .

These values represent averages and FWHM of gaussians adjusted to each histogram of abundance distribution. It should be noted that nitrogen elemental abundances depend on OII ionic abundances used as icf, so uncertainties in this last parameter will increase the dispersion found in N. As nitrogen is also a product of stellar nucleosynthesis dredged-up during the stellar evolution of the progenitors, their mass spectrum will also contribute to the derived dispersion in N abundance.

Sulphur abundances have to be taken cautiously because they were derived from relatively weak [SII] λ 6716+31 and [SIII] λ 6312 lines. Besides, sulphur abundances in particular are strongly dependent on electron temperatures.

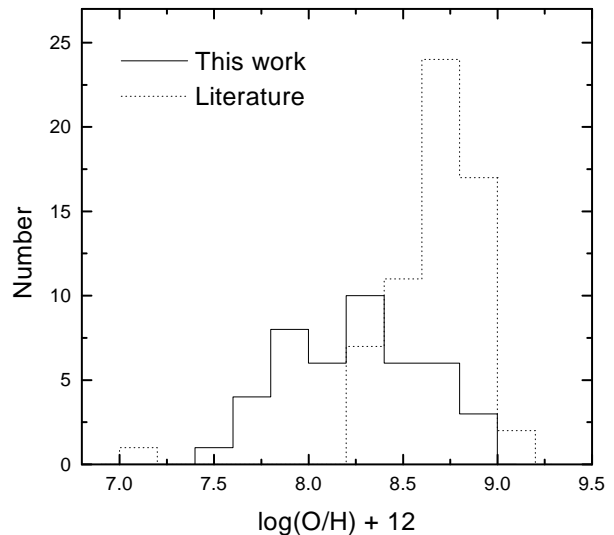


Fig. 4. Distribution of oxygen abundances for our sample (full line) and for the literature (dashed line).

Figures 4 and 5 display the abundance distribution for our sample compared to that from other works, respectively for oxygen and $\log(N/O)$. At a first glance, distributions seem to be different, but this is due to different

Table 5. Ionic abundances

Name	HeI	HeII	NII	SII	SIII	OII	OIII	ArIII
K 5-1	0.108	-	9.29E-07	2.69E-08	4.57E-07	2.47E-06	7.29E-05	5.76E-07
K 5-3	0.082	0.042	4.19E-06	1.82E-07	9.44E-07	3.16E-05	1.17E-04	4.06E-07
K 5-4	0.109	-	1.02E-05	3.31E-07	3.97E-06	5.34E-05	3.91E-04	1.12E-06
K 5-5	0.03	-	5.63E-07	4.62E-08	2.93E-07	2.06E-06	5.08E-05	2.02E-07
K 5-6	0.023	0.09	4.06E-07	4.16E-08	5.59E-07	2.22E-06	7.07E-05	2.09E-07
K 5-7	0.066	0.094	1.25E-06	7.38E-08	-	5.22E-06	8.44E-05	6.10E-07
K 5-9	0.142	0.015	7.73E-06	2.50E-07	1.29E-06	4.17E-05	2.32E-04	1.14E-06
K 5-11	0.174	0.017	1.13E-04	2.36E-06	3.19E-06	4.57E-04	2.22E-04	1.64E-06
K 5-12	0.046	0.06	2.91E-06	2.05E-07	3.11E-06	1.25E-05	3.45E-04	5.61E-06
K 5-13	0.082	0.058	8.68E-05	2.31E-06	4.20E-06	1.75E-04	1.95E-04	1.53E-06
K 5-14	0.094	0.022	2.50E-05	8.56E-07	1.68E-06	8.88E-05	1.71E-04	9.26E-07
K 5-16	0.129	0.012	1.28E-04	4.13E-06	1.78E-05	4.12E-04	4.47E-04	2.32E-06
K 5-17	0.076	-	1.77E-06	1.73E-07	6.41E-07	6.97E-06	1.52E-04	3.66E-07
K 5-19	0.071	0.082	3.34E-06	1.73E-07	-	7.57E-06	5.29E-05	2.30E-07
K 5-20	0.128	-	7.30E-07	3.08E-08	7.78E-07	3.17E-05	2.14E-04	4.74E-07
SB 01	0.046	0.032	1.79E-06	1.68E-07	-	1.48E-05	8.26E-05	1.70E-07
SB 02	0.141	0.027	6.96E-06	2.48E-07	-	1.99E-05	5.90E-05	3.62E-07
SB 03	0.105	0.054	1.59E-06	6.55E-08	-	1.24E-06	3.09E-05	1.78E-07
SB 04	0.095	0.019	6.34E-06	1.75E-07	-	6.12E-06	3.86E-05	1.85E-07
SB 06	0.084	0.046	2.18E-05	1.27E-07	-	1.22E-05	2.65E-05	3.72E-07
SB 12	0.049	0.066	2.93E-07	-	-	-	4.08E-05	2.52E-07
SB 15	0.053	0.046	7.64E-08	6.16E-08	-	4.78E-06	7.79E-05	4.21E-08
SB 17	0.142	-	3.06E-05	1.45E-06	-	8.98E-05	2.58E-06	5.47E-08
SB 18	0.046	0.057	1.98E-07	2.25E-08	-	6.93E-06	7.53E-05	6.85E-08
SB 19	0.156	-	2.09E-06	2.48E-07	5.05E-06	3.16E-05	3.34E-04	1.85E-06
SB 20	0.116	0.028	5.90E-08	2.64E-08	-	1.37E-06	6.02E-05	4.48E-08
SB 21	0.021	0.102	1.02E-07	-	-	3.11E-06	5.04E-05	2.01E-07
SB 24	0.032	0.077	5.68E-07	1.56E-07	-	1.67E-05	4.79E-05	1.62E-07
SB 25	0.138	-	3.24E-05	9.02E-07	8.16E-07	2.94E-05	8.54E-05	8.10E-07
SB 26	0.016	0.082	2.19E-07	-	1.45E-07	1.20E-06	8.96E-05	1.50E-07
SB 28	0.004	0.11	3.86E-08	-	-	-	1.84E-05	5.70E-08
SB 30	0.061	0.04	1.92E-06	6.88E-08	3.46E-07	1.22E-05	1.27E-04	1.87E-07
SB 31	0.009	0.101	3.33E-07	2.17E-08	8.03E-07	2.27E-06	6.84E-05	1.59E-07
SB 32	0.12	0.022	6.71E-06	4.40E-07	2.66E-06	2.43E-05	4.32E-05	2.05E-07
SB 33	0.079	0.016	7.51E-06	3.55E-07	-	1.04E-05	2.43E-05	2.85E-07
SB 34	0.01	0.129	4.09E-07	-	-	-	2.37E-05	7.48E-08
SB 35	0.107	-	6.09E-07	4.24E-08	-	1.04E-05	1.67E-04	3.87E-07
SB 37	-	0.146	7.93E-07	4.91E-08	1.13E-06	1.06E-06	9.73E-05	-
SB 38	0.077	0.058	1.24E-07	3.83E-08	-	4.76E-07	6.68E-05	1.40E-07
SB 42	0.051	0.004	1.08E-06	4.22E-08	-	1.39E-05	4.70E-05	1.52E-07
SB 44	0.029	0.107	3.25E-06	1.39E-07	-	5.30E-05	1.27E-04	6.99E-07
SB 50	0.056	0.086	7.41E-07	8.29E-08	-	3.56E-06	5.02E-05	2.73E-07
SB 52	0.033	0.103	3.93E-07	5.52E-08	6.56E-07	7.37E-06	8.44E-05	4.09E-07
SB 53	0.015	0.119	5.56E-07	3.81E-08	1.45E-06	1.10E-06	7.61E-05	4.75E-07
SB 55	0.035	0.118	1.08E-07	3.64E-09	-	2.02E-06	1.50E-05	8.31E-08

coverages in the sky. Our sample spans over 20 deg. in galactic longitude and 11 deg. in latitude, in a nearly homogeneous distribution, while data from the literature are strongly concentrated nearer the position of the galactic center. These different distributions can be seen in fig. 6, where both samples are mapped in galactic co-ordinates.

N/O distribution reflects the mass distribution of the progenitors, with more massive objects having higher N

abundances. A few objects with $\log(N/O) < -1$ could be originated from old, small-mass progenitors. The distribution peaks for typical type-II PNe abundances.

Concerning abundances, fig. 7 displays the oxygen abundance distribution in latitude for our sample, compared with results from the literature taken from the three works mentioned before. Examining the results, we see in our sample a concentration of objects around $\varepsilon(O)=8.0$ at

Table 6. Elemental abundances

Planetary	He	σ_{He}	$\varepsilon(N)$	σ_N	$\varepsilon(S)$	σ_S	$\varepsilon(O)$	σ_O	$\varepsilon(Ar)$	σ_{Ar}	Notes
K 5-1	0.108	0.031	7.45	0.30	6.91	0.68	7.88	0.21	5.9	0.16	
K 5-3	0.123	0.019	7.48	0.20	6.45	0.27	8.35	0.16	6.02	0.17	
K 5-4	0.109	0.014	7.93	0.17	7.24	0.28	8.65	0.13	6.23	0.14	
K 5-5	0.030	0.013	7.16	0.24	6.67	0.48	7.72	0.19	5.45	0.17	1
K 5-6	0.114	0.017	7.81	0.21	7.05	0.18	8.55	0.14	6.15	0.15	
K 5-7	0.160	0.031	7.72	0.14	5.8	0.17	8.33	0.17	6.32	0.18	
K 5-9	0.157	0.030	7.75	0.15	6.7	0.31	8.48	0.34	6.3	0.26	
K 5-11	0.191	0.056	8.26	0.22	6.92	0.38	8.87	0.45	6.87	0.67	
K 5-12	0.106	0.013	8.28	0.28	7.71	0.51	8.91	0.09	7.25	0.10	
K 5-13	0.140	0.029	8.5	0.11	7.03	0.20	8.8	0.20	6.82	0.27	
K 5-14	0.116	0.019	7.95	0.11	6.68	0.15	8.5	0.14	6.37	0.22	
K 5-16	0.141	0.042	8.47	0.09	7.56	0.31	8.97	0.33	6.82	0.64	
K 5-17	0.076	0.015	7.61	0.17	6.99	0.25	8.2	0.11	5.71	0.12	
K 5-19	0.153	0.031	7.76	0.25	5.82	0.20	8.11	0.16	5.88	0.22	
K 5-20	0.128	0.027	6.75	0.13	6.48	0.21	8.39	0.18	5.86	0.19	
SB 01	0.077	0.016	7.3	0.16	5.74	0.16	8.22	0.23	5.66	0.19	1
SB 02	0.168	0.024	7.52	0.13	5.74	0.09	7.97	0.11	5.89	0.12	
SB 03	0.159	0.020	7.8	0.11	5.96	0.1	7.69	0.08	5.58	0.09	2,3
SB 04	0.114	0.028	7.75	0.12	5.79	0.11	7.73	0.13	5.54	0.13	
SB 06	0.130	0.029	8.03	0.09	5.4	0.13	7.78	0.17	6.05	0.15	
SB 12	0.115	0.034	-	-	-	-	7.98	0.18	5.9	0.25	
SB 15	0.099	0.017	6.39	0.17	5.72	0.13	8.19	0.07	5.05	0.19	
SB 17	0.142	0.033	7.5	0.15	6.32	0.16	7.97	0.34	6.42	0.35	
SB 18	0.103	0.023	6.72	0.26	5.11	0.23	8.26	0.15	5.35	0.21	
SB 19	0.156	0.031	7.38	0.19	7.47	0.28	8.56	0.21	6.43	0.16	3
SB 20	0.144	0.029	6.52	0.17	5.85	0.18	7.88	0.15	4.88	0.19	3
SB 21	0.123	0.017	7.02	0.16	-	-	8.5	0.11	6.23	0.13	
SB 24	0.109	0.018	6.87	0.30	5.53	0.37	8.34	0.19	6	0.21	
SB 25	0.138	0.038	8.1	0.15	6.58	0.22	8.06	0.25	6.16	0.21	3
SB 26	0.098	0.013	8	0.17	6.88	0.26	8.74	0.10	6.09	0.12	
SB 28	0.114	0.021	-	-	-	-	8.75	0.27	6.37	0.28	
SB 30	0.100	0.010	7.56	0.12	6.35	0.17	8.36	0.06	5.66	0.06	
SB 31	0.110	0.011	8.13	0.20	7.15	0.19	8.96	0.15	6.45	0.14	
SB 32	0.142	0.032	7.34	0.14	6.76	0.20	7.9	0.15	5.71	0.16	
SB 33	0.095	0.029	7.48	0.13	5.85	0.12	7.62	0.16	5.82	0.16	
SB 34	0.139	0.023	-	-	-	-	8.53	0.17	6.16	0.19	
SB 35	0.107	0.016	7.02	0.12	5.55	0.15	8.25	0.10	5.74	0.11	
SB 37	0.146	0.015	-	-	7.91	0.17	-	-	-	-	
SB 38	0.135	0.017	7.49	0.16	6.65	0.19	8.07	0.07	5.52	0.12	3
SB 42	0.055	0.010	6.71	0.19	5	0.10	7.82	0.11	5.45	0.11	
SB 44	0.136	0.017	7.72	0.18	5.45	0.17	8.93	0.19	6.8	0.18	
SB 50	0.141	0.021	7.45	0.15	5.79	0.17	8.13	0.12	6	0.12	3
SB 52	0.137	0.011	7.31	0.34	6.63	0.59	8.58	0.09	6.39	0.10	
SB 53	0.134	0.012	8.53	0.2	7.85	0.4	8.83	0.14	6.75	0.12	3
SB 55	0.153	0.026	6.6	0.32	4.17	0.40	7.87	0.19	5.74	0.24	

1 Spectra with large errors

2 An average temperature was used

3 Lines [OII] λ 3726+29 were used to derive OII abundance

latitudes greater than 5 deg., which deviates from the average distribution. These objects belong to a region where no PNe abundances have been determined before. They can be thick disk objects seen towards the bulge, or effec-

tively bulge objects at higher altitudes above the galactic plane. If their membership is confirmed, their lower abundances suggest the existence of a vertical abundance gradient within the bulge.

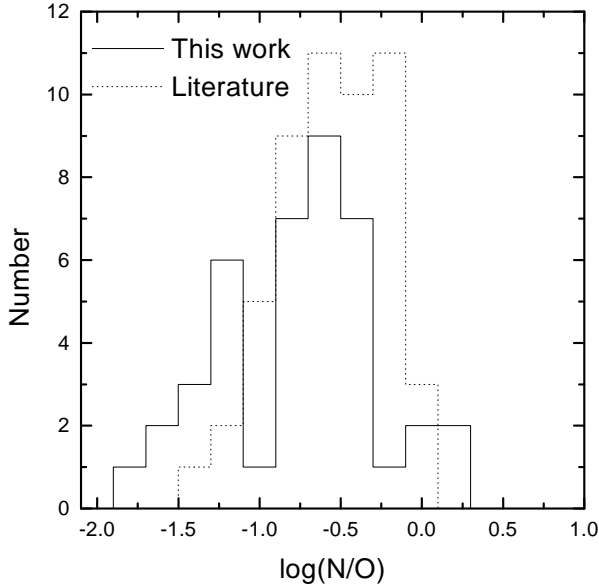


Fig. 5. Distribution of $\log(N/O)$ for our sample (full line) and for the literature (dashed line).

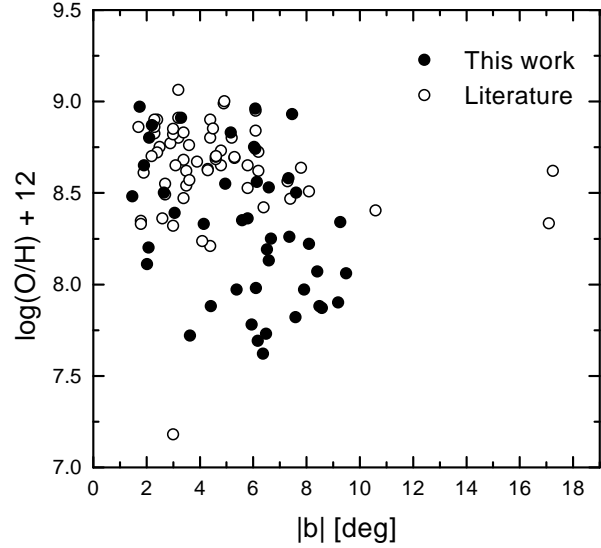


Fig. 7. Oxygen abundances vs. absolute value of galactic latitude for our sample (filled circles) and for the literature (open circles).

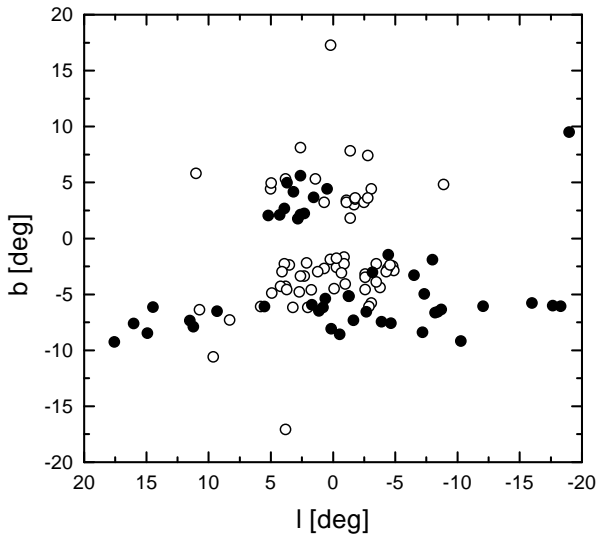


Fig. 6. Space distribution for our sample (filled circles) and for the literature (open circles).

Figure 8 shows $\log(N/O)$ versus helium abundance for our sample and for the same literature objects used above. As nitrogen and helium are nucleosynthesis products and related to the mass spectrum of the progenitors, they are expected to be correlated. In the figure there is an upper group of objects that, in spite of the expected dispersion due to the mass distribution of the progenitors, shows a

correlation. It can also be seen that there is another small group with low N/O and He/H varying from 0.10 to 0.16. The same pattern appears in the results of Cuisinier et al. (1996) and can be related to objects at high z above the galactic plane. However, as distances for our these object are highly uncertain, this is still an open question.

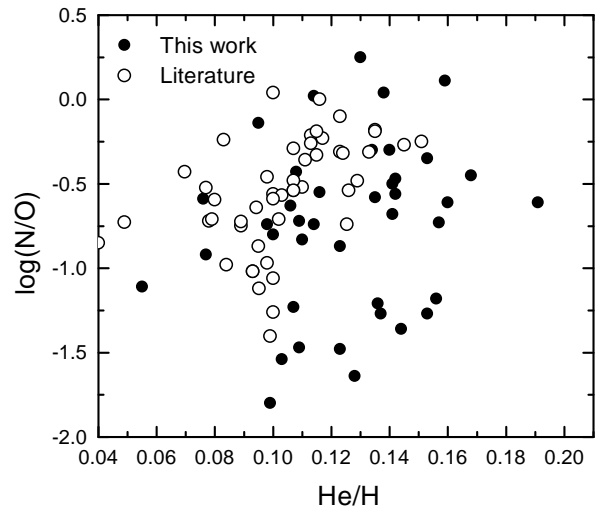


Fig. 8. $\log(N/O)$ vs. helium abundance for our sample (filled circles) and for the literature (open circles).

One has to take into account that bulge membership is an additional problem hard to solve for PNe. Objects from Beaulieu et al. (1999) are distributed over a wide range of galactic latitudes and their elemental abundances are lower than those from other works. This problem will be tackled in another work, now in preparation.

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